

# Outflows and Jets: Theory and Observations

Summer term 2011

Henrik Beuther & Christian Fendt

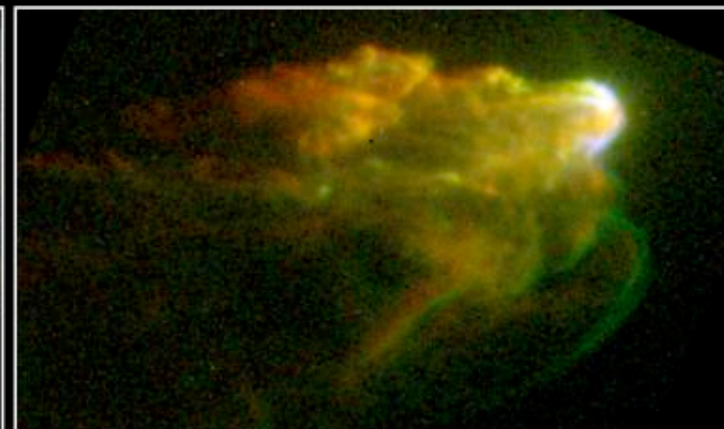
## **15.04 Today: Introduction & Overview (H.B. & C.F.)**

- 29.04 Definitions, parameters, basic observations (H.B.)
- 06.05 Basic theoretical concepts & models (C.F.)
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- 22.07 Theory of relativistic jets (C.F.) (examination week?)

More Information and the current lecture files: [http://www.mpia.de/homes/beuther/lecture\\_ss11.html](http://www.mpia.de/homes/beuther/lecture_ss11.html)  
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# Discovery of outflows I

Herbig 1950, 1951; Haro 1952, 1953



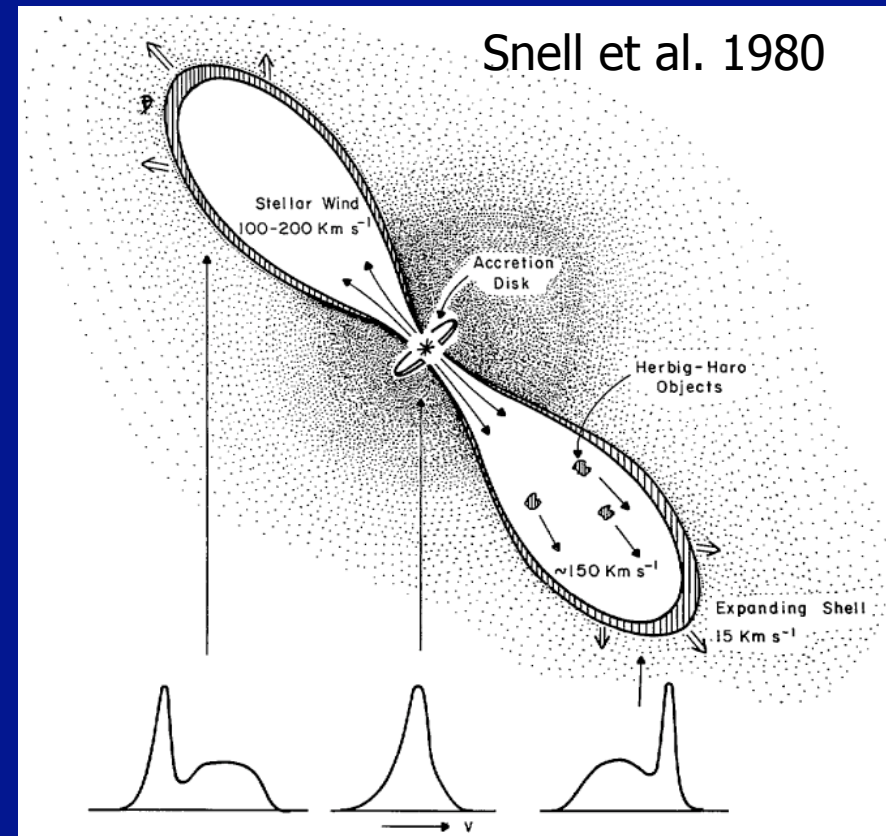
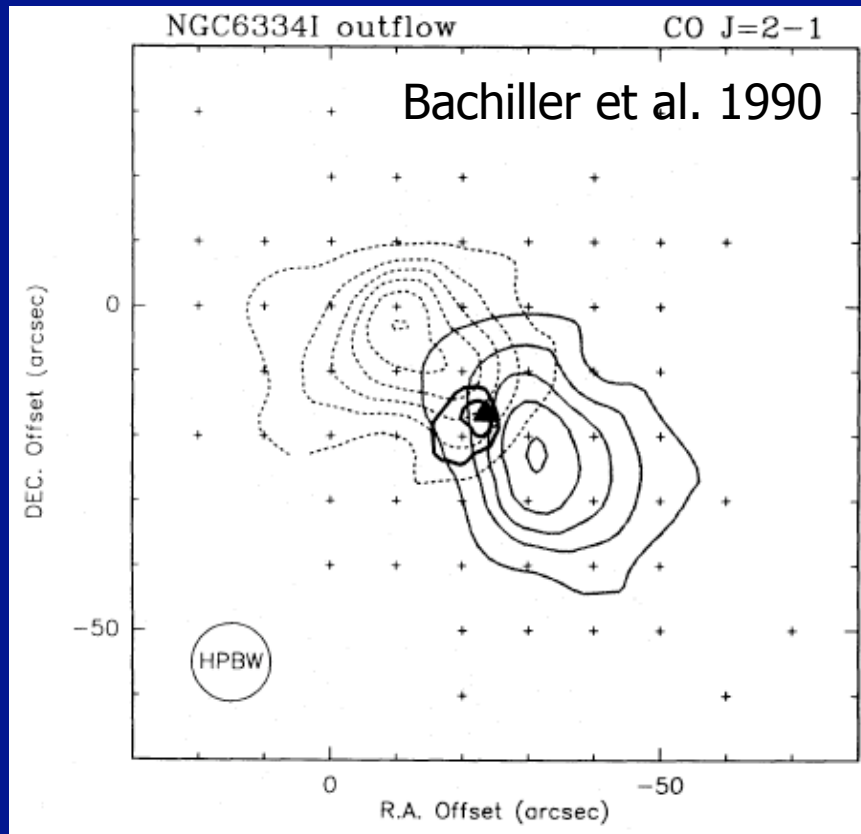
**Jets from Young Stars · HH1/HH2**

**HST · WFPC2**

PRC95-24c · ST ScI OPO · June 6, 1995 · J. Hester (AZ State U.), NASA

Initially thought to be embedded protostars but soon spectra were recognized as caused by shock waves → jets and outflows

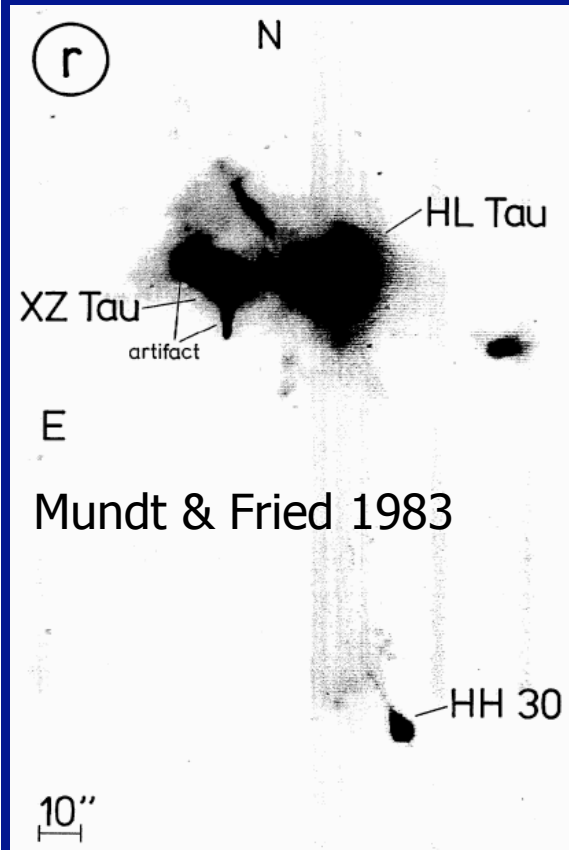
# Discovery of outflows II



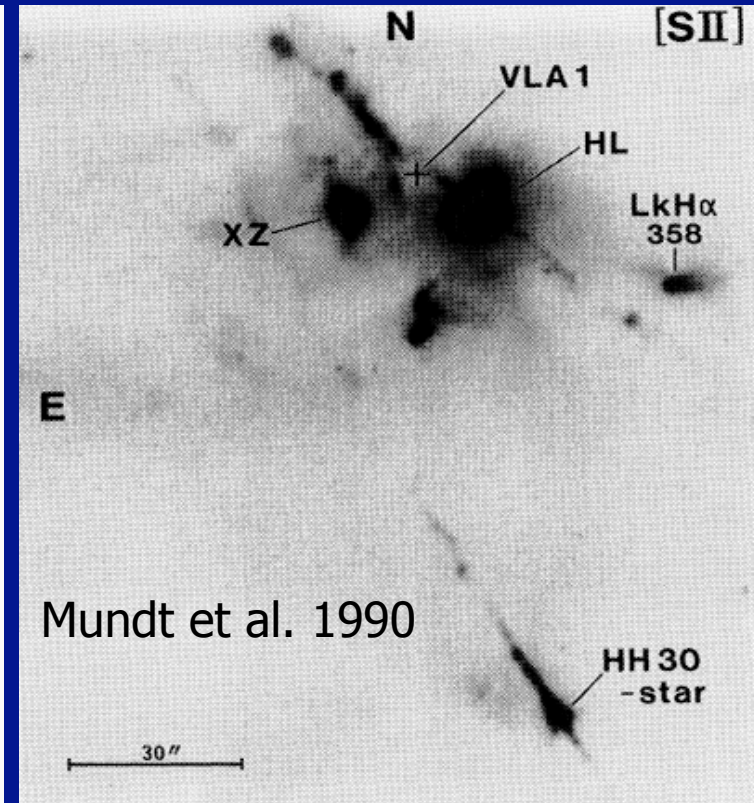
- In the mid to late 70th, first CO non-Gaussian line wing emission detected (Kwan & Scoville 1976).
- Bipolar structures, extremely energetic, often associated with HH objects

# Optical jet observations

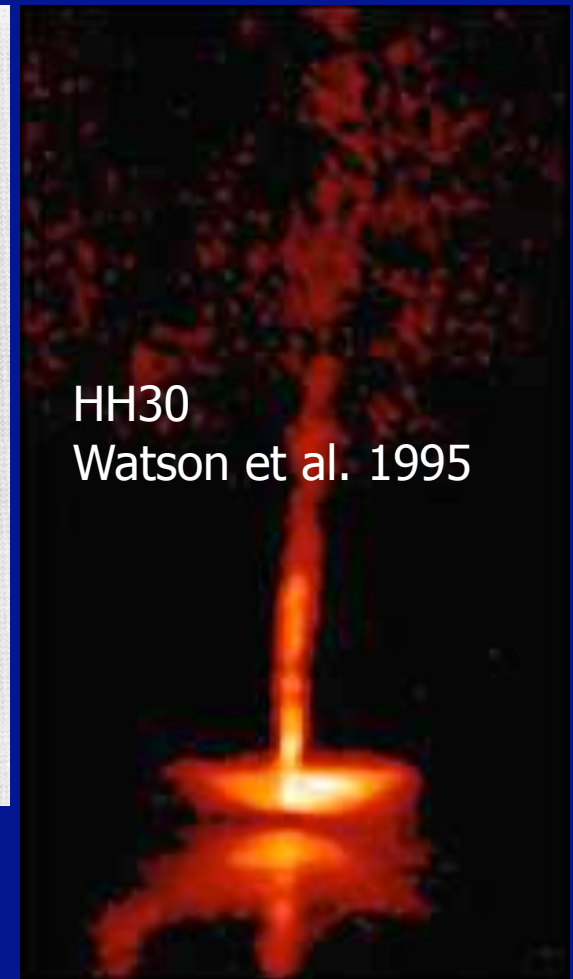
1983



1990 (both ground based)



1995 (HST)

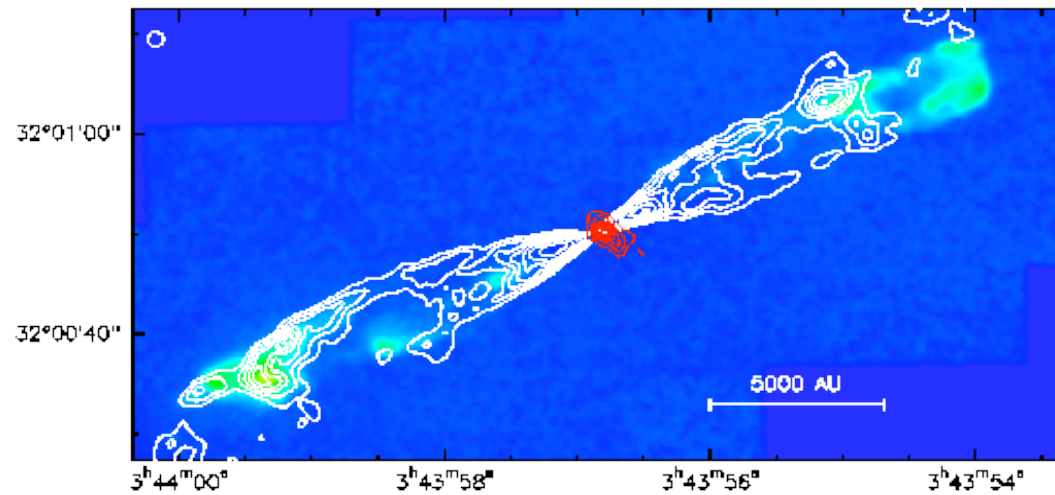




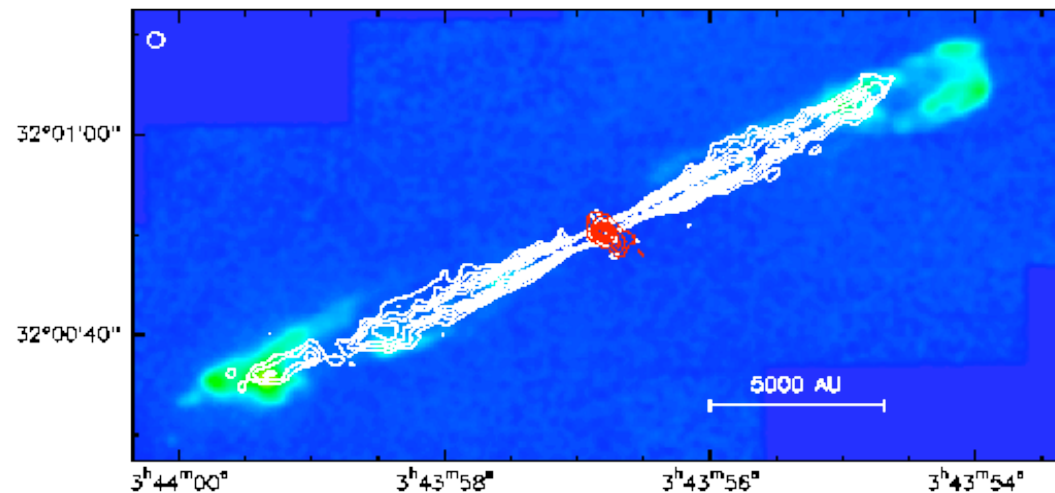
# The prototypical molecular outflow HH211

HH211, Gueth et al. 1999

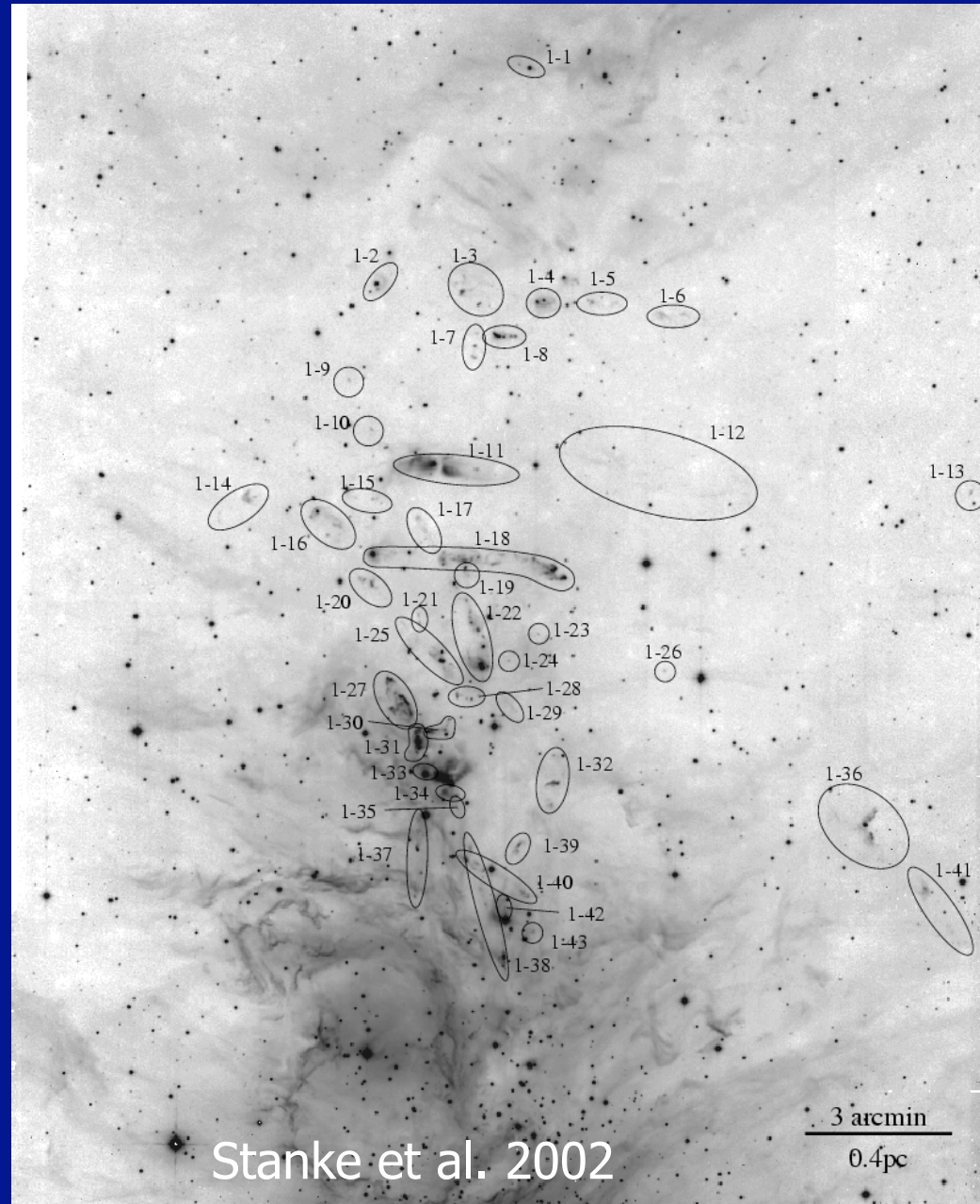
H<sub>2</sub> 2.12 μm (colors) + CO J=2-1 v<10 km/s (white) + continuum 1.3 mm (red)



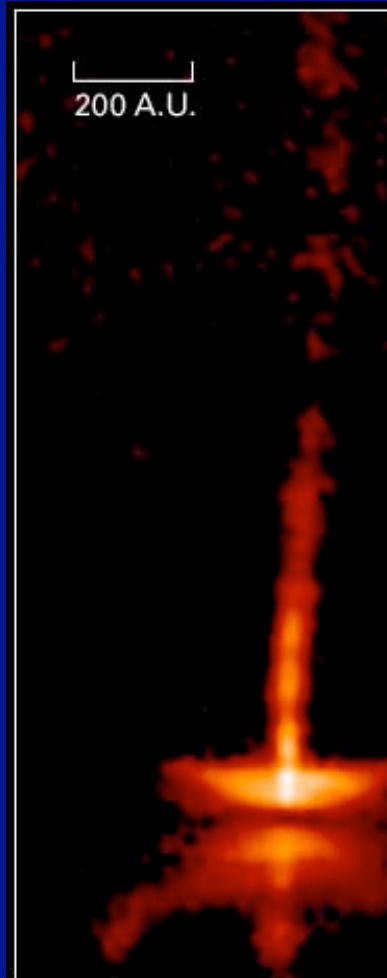
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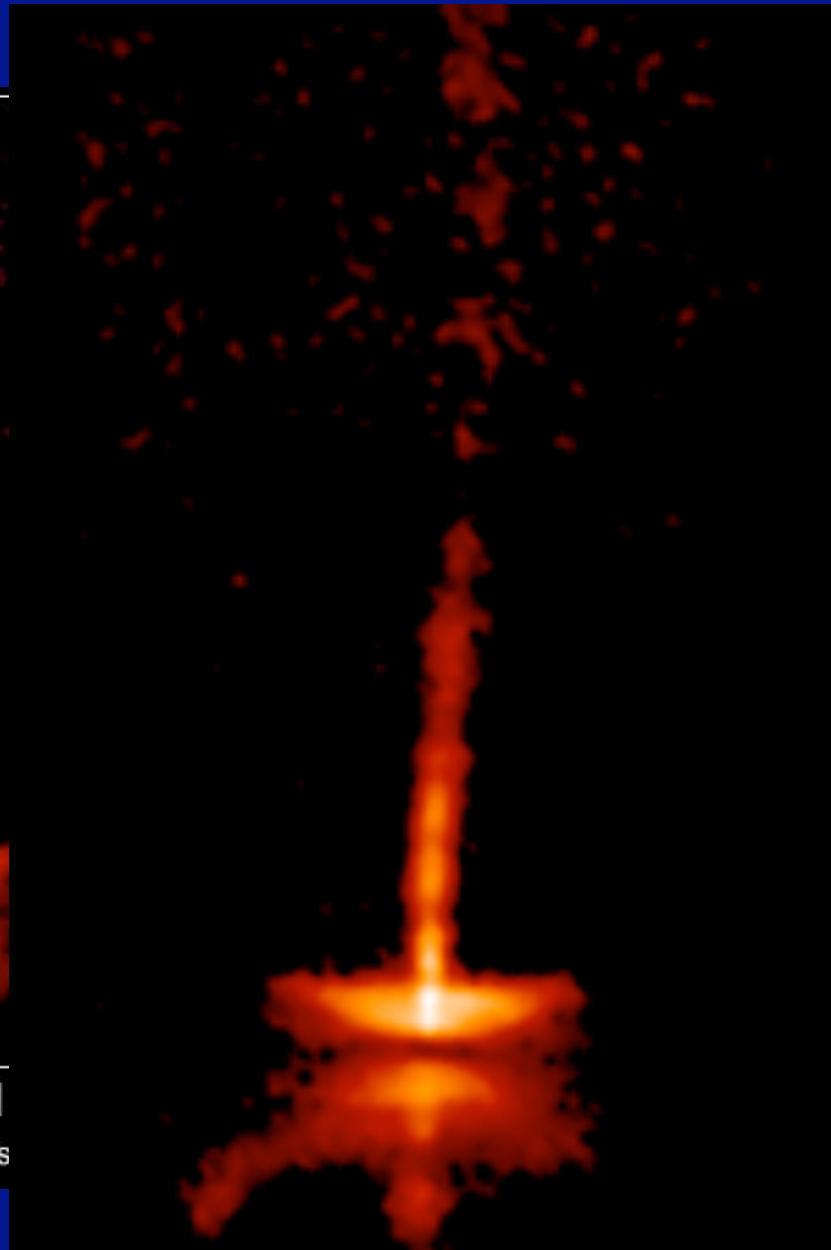
# Outflow multiplicities in Orion



# HH30, a disk-outflow system



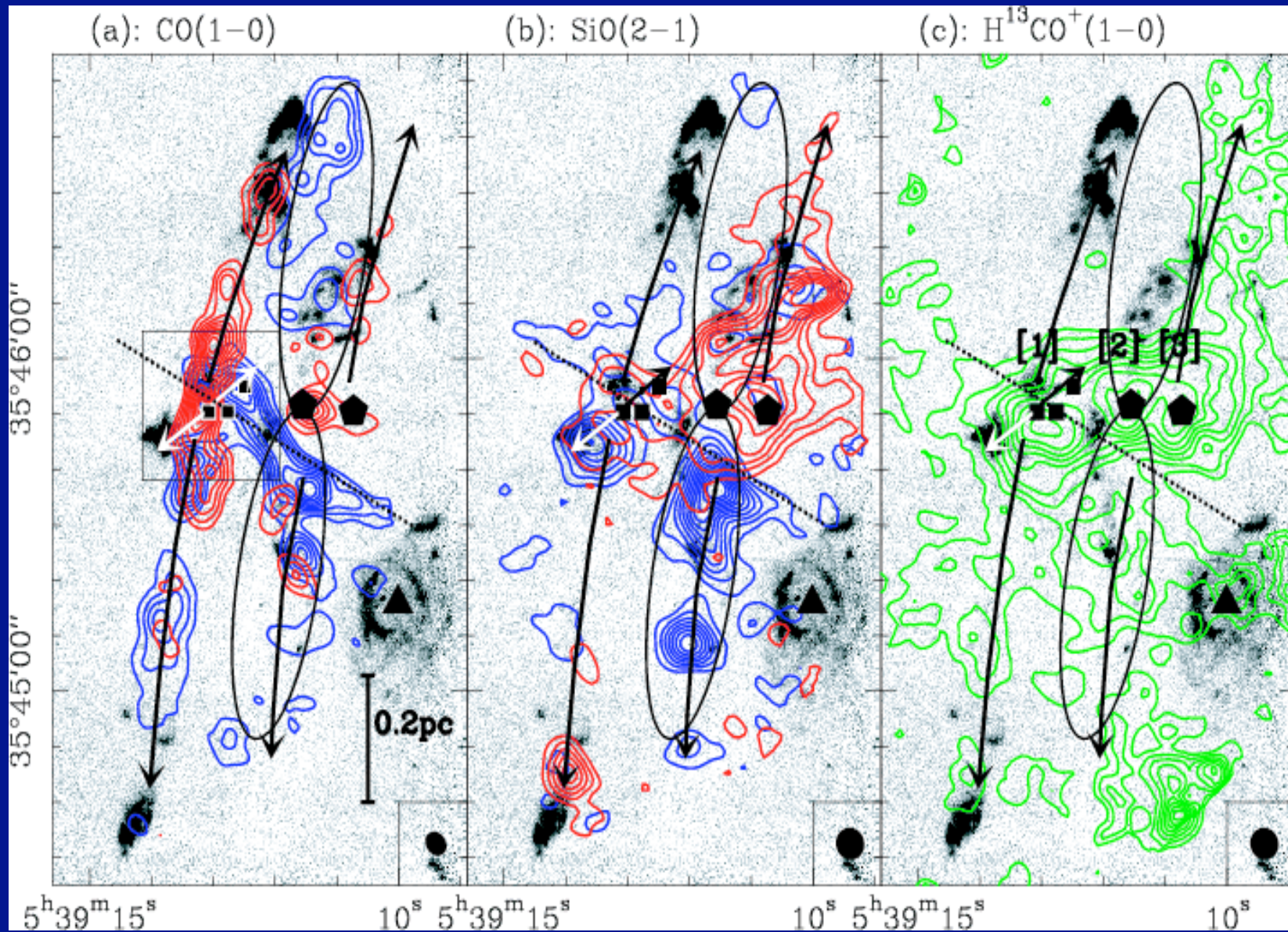
**The Dynamic HH**  
NASA and A. Watson (Ins



**HST • WFPC2**  
2b



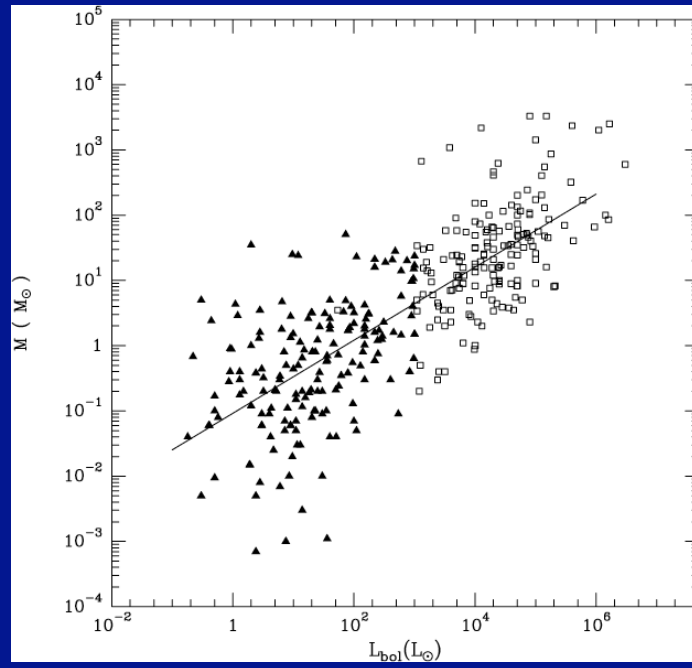
# Outflows from massive star-forming regions



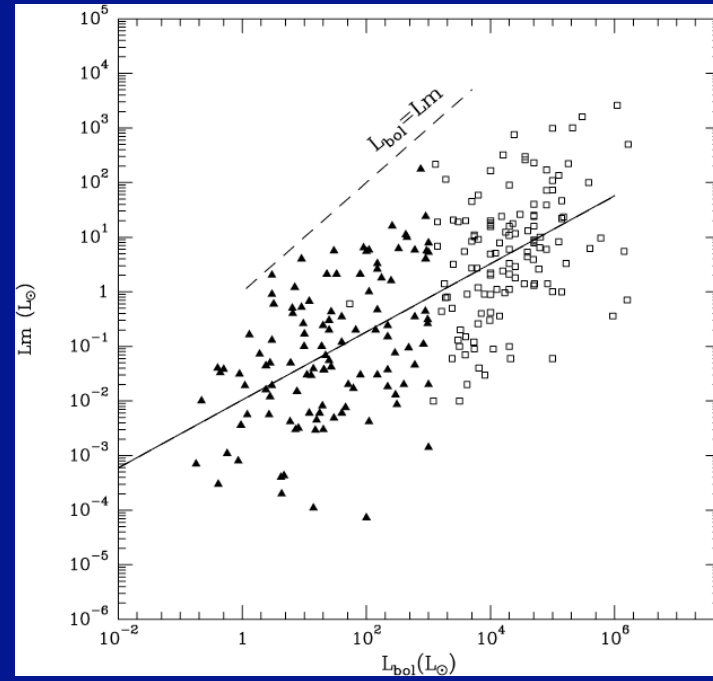


# Outflow properties

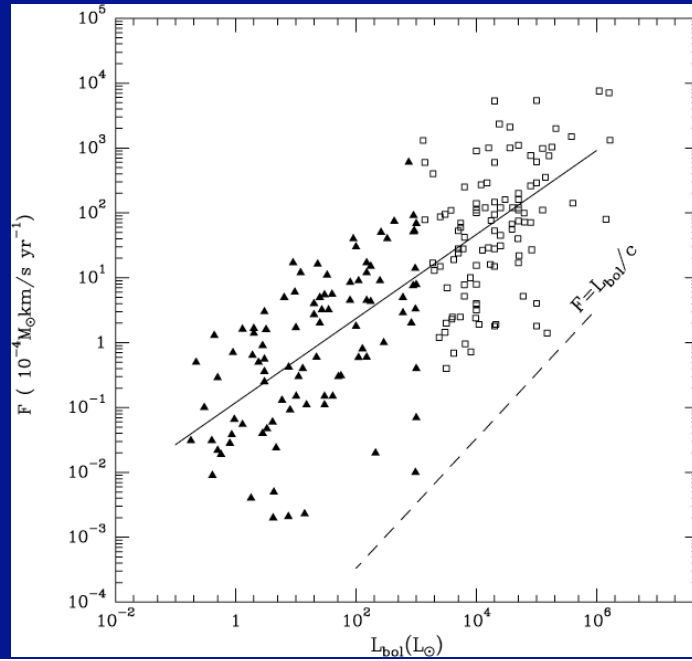
Mass



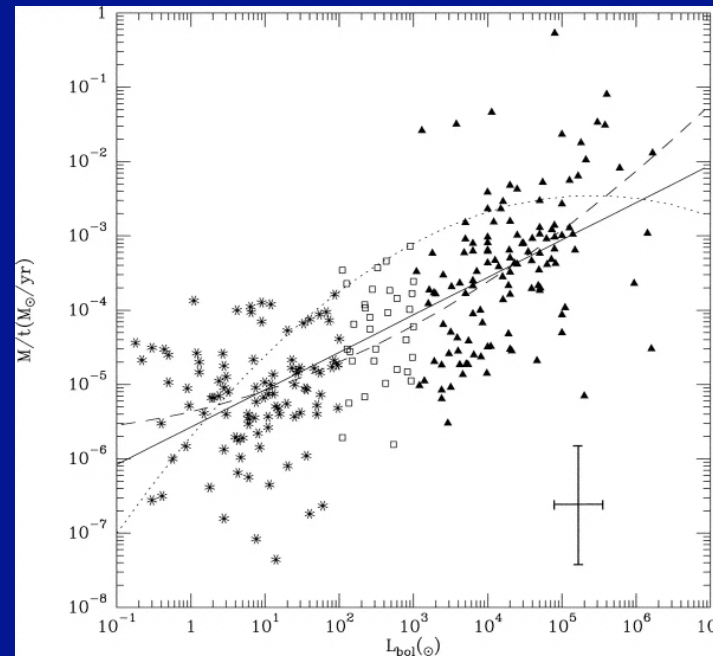
Mechanical L



Mechanical Force



Outflow rate

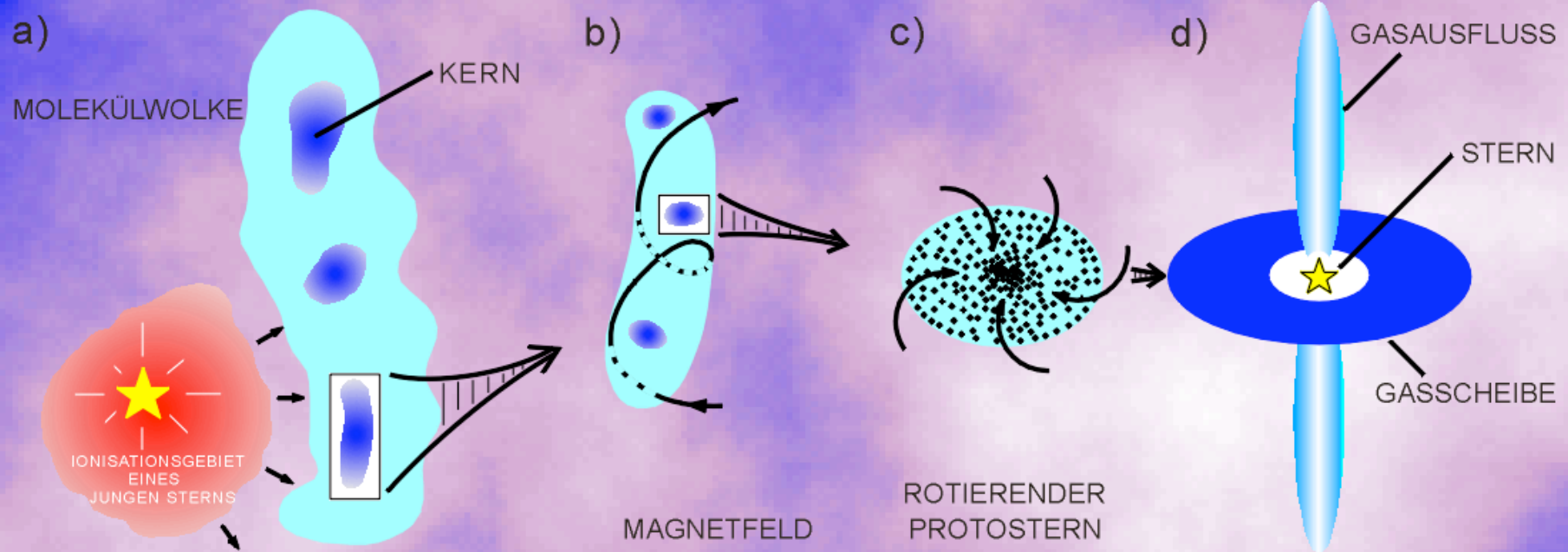


# General jet/outflow parameters

- Initial jet-velocities  $> 150\text{km/s}$ , can be a few  $100\text{km/s}$  jets even  $>500\text{km/s}$
- Entrained outflow gas velocities usually a few  $10\text{km/s}$
  
- Jets extremely collimated with opening angles of a few degrees, molecular outflows can be less collimated with opening angles up to  $90$  degrees.
- Length between  $1000\text{AU}$  and a couple of parsec
  
- Outflow masses between sub-solar to  $\sim 100M_{\text{sun}}$ .
- Outflow rates from a few times  $10^{-8}M_{\text{sun}}/\text{yr}$  for low-mass T Tauri stars to a few times  $10^{-3}M_{\text{sun}}/\text{yr}$  for high-mass star-forming regions.
- Outflow luminosities between  $10^{-3}$  and several  $100 L_{\text{sun}}$
  
- Jets largely neutral with ionization degrees between  $10^{-1}$  and  $10^{-8}$ .
- Magnetic fields are observed and jets are MHD driven.
- Momentum conservation between primary jet and entrained outflow gas.

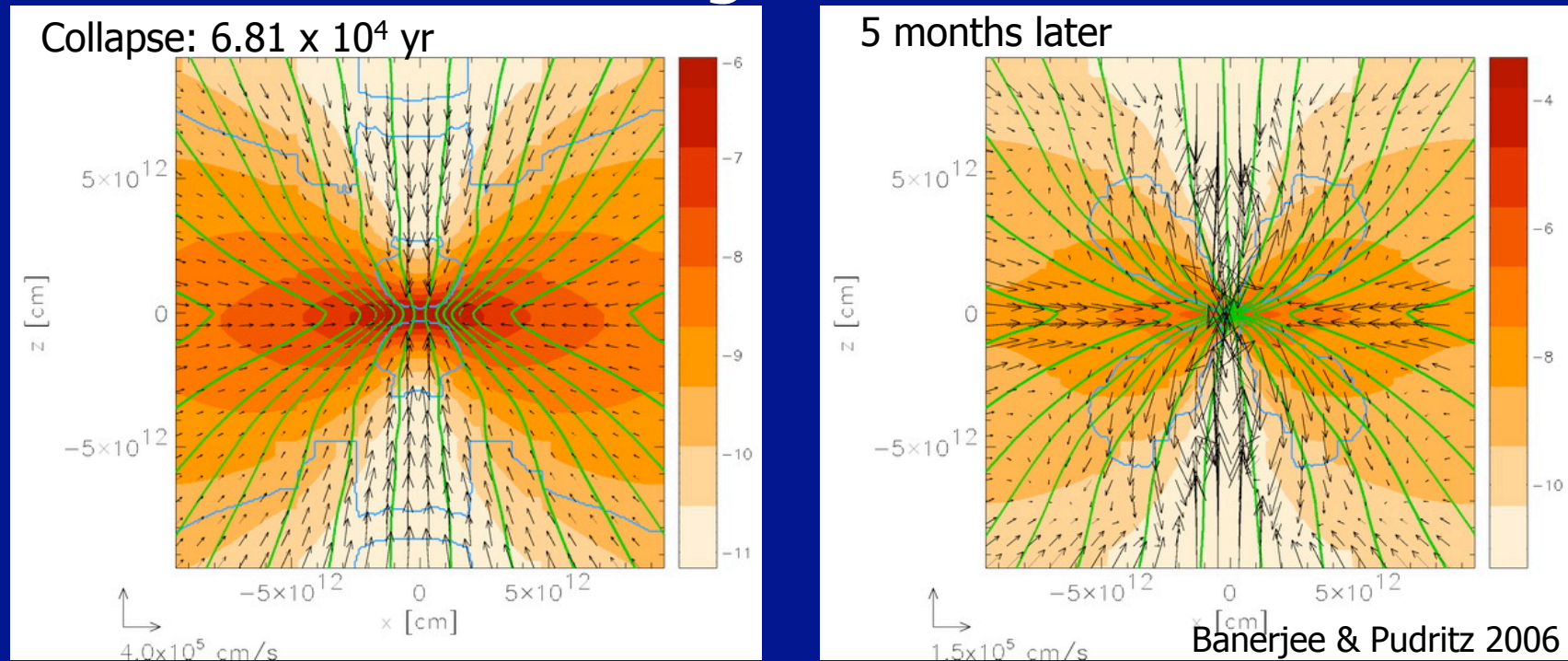
# Star Formation Paradigm

## DIE ENTWICKLUNGSSTUFEN DER STERNENTSTEHUNG



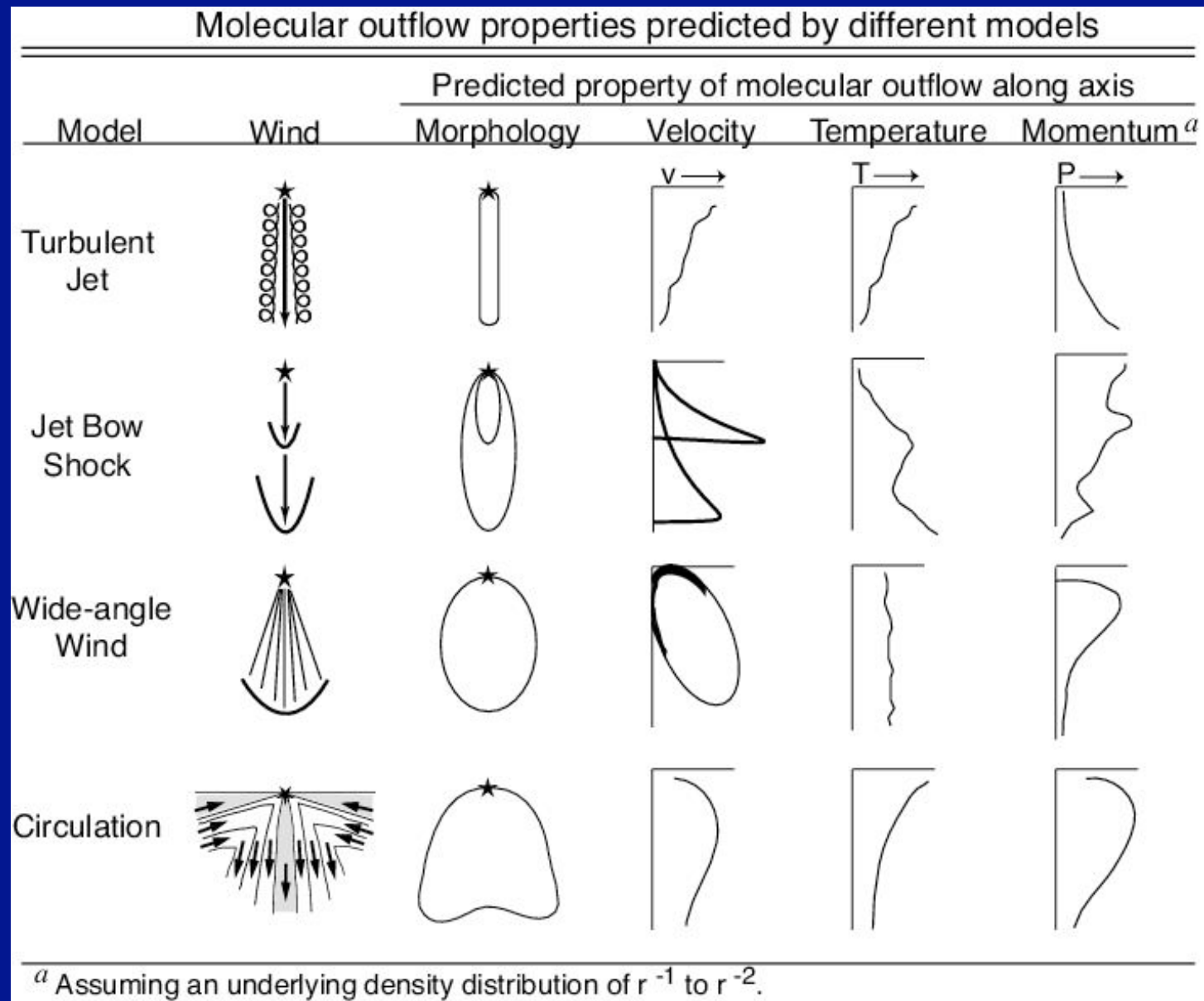


# Jet-launching: Disk winds



- Infalling core pinches magnetic field.
- If poloidal magnetic field component has angle larger  $30^\circ$  from vertical, centrifugal forces can launch matter-loaded wind along field lines from disk surface.
- Wind transports away from 60 to 100% of disk angular momentum.

# Outflow entrainment models



# Why studying outflows/jets?

Interesting astrophysical phenomena in itself.

Important for angular momentum removal.

Impact on ISM.

Driving Turbulence.

Potentially triggering other star formation.

Shock chemistry.

...



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