Galaxy Dynamics

Galaxies Block Course 2008
Hans-Walter Rix - MPIA

For more detail on any of this

Binney and Tremaine, 1996

Binney and Merrifield, 2002

Basic Goals of Stellar Dynamics

- Most Empirical:
  - What is the (total) mass distribution?
  - On what orbits do the constituent mass elements, or the tracer masses, move?

- Going a bit deeper:
  - What are the mass components: stars, gas, dark matter, etc…
  - Is the mass budget accounted for by the known/identified mass constituents?
  - Are (most) systems in (approximate) steady state on a "dynamical timescale".

- Finally:
  - is the range of observed galaxy structures determined by stability?
  - What can be learned about the formation process of galaxies from their dynamical state?
  - Any slow ("secular") internal re-shaping?

•But, let's not forget the practical question:

How do we use observable information to get these answers?

Observables:

•Spatial distribution and kinematics of "tracer population(s)", which may make up

- •all (in globular clusters?)
- •much (stars in elliptical galaxies?) or
- •little (ionized gas in spiral galaxies)

of the "dynamical" mass.

In external galaxies only 3 of the 6 phase-space dimensions, $x_{proj}, y_{proj}, v_{LOS}$, are observable!

Note: since $t_{dynamical} \sim 10^8$ yrs in galaxies, observations constitute an instantaneous snapshot.

...the Galactic Center is an exciting exception..

Gas vs. Stars
or
Collisionless vs. Collisional Matter

How often do stars in a galaxy „collide"?

- $R_{Sun} \approx 7 \times 10^{10}$ cm; $D_{Sun-\alpha Cen} \approx 10^{19}$ cm!
  => collisions extremely unlikely!

  …and in galaxy centers?

  Mean surface brightness of the Sun is μ =-11mag/sqasec, which is distance independent. The central parts of other galaxies have μ~ 12 mag/sqasec. Therefore, $(1 - 10^{-9})$ of the projected area is empty.

⇒ Even near galaxy centers, the path ahead of stars is empty.

Dynamical time-scale (=typical orbital period)

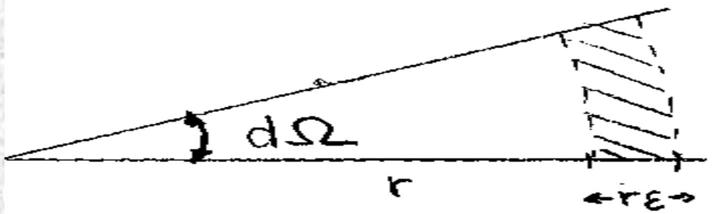
Milky Way:

R~8kpc v~200km/s → t_{orb}~240 Myrs

→ $t_{orb} \sim t_{Hubble}/50$ ← true for galaxies of most scales

Stars in a galaxy feel the gravitational force of other stars. But of which ones?

- consider homogeneous distribution of stars, and force exerted on one star by other stars seen in a direction $d\Omega$ within a slice of $[r, r\times(1+\varepsilon)]$



=> $dF \sim GdM/r^2 = G\rho \times r(!) \times \varepsilon d\Omega$

- gravity from the multitude of distant stars dominates!

What about (diffuse) interstellar gas?

- continuous mass distribution
- gas has the ability to lose (internal) energy through radiation.
- Two basic regimes for gas in a potential well of ‚typical orbital velocity', v
  - $kT/m \approx v^2 \rightarrow$ hydrostatic equilibrium
  - $kT/m \ll v^2$, as for atomic gas in galaxies
- in the second case:

  supersonic collisions → shocks → (mechanical) heating → (radiative) cooling → energy loss

  For a given (total) angular momentum, what's the minimum energy orbit?

  A (set of) concentric (co-planar), circular orbits.

  => cooling gas makes disks!

I. Describing Stellar Dynamical Systems in Equilibrium

Modeling Collisionless Matter: Approach I

Phase space: d$\underline{x}$, d$\underline{v}$

We describe a many-particle system by its distribution function f(x,v,t) = density of stars (particles) within a phase space element

Starting point: Boltzmann Equation (= phase space continuity equation)
It says: if I follow a particle on its gravitational path (=Lagrangian derivative) through phase space, it will always be there.

$$\frac{Df(\vec{x},\vec{v},t)}{Dt} = \frac{\partial f}{\partial t} + \vec{v}\,\frac{\partial f}{\partial \vec{x}} - \vec{\nabla}\Phi_{grav}\,\frac{\partial f}{\partial \vec{v}} = 0$$

A rather ugly partial differential equation!

Note: we have substituted gravitational force for accelaration!

To simplify it, one takes velocity moments:
i.e. $\int_{\mathbb{R}^3} \ldots v^n d^3v$ n = 0,1, ... on both sides

Moments of the Boltzmann Equation

0th Moment

$$\frac{\partial \rho}{\partial t} + \vec{\nabla} \cdot (\rho \vec{u}) = 0$$

mass conservation

ρ: mass density; v/u: indiv/mean particle velocity

1st Moment $\int \ldots v_j d^3v$

$$\frac{\partial}{\partial t}(\rho \vec{u}) + \vec{\nabla} \cdot \left(\rho \left(\underline{\underline{T}} + \vec{u} \cdot \vec{u}\right)\right) + \rho \vec{\nabla} \Phi = 0$$

$$\text{with } \rho \underline{\underline{T}} = \int f \cdot (v_i - u_i)(v_j - u_j)\, d^3v$$

"Jeans Equation"

The three terms can be interpreted as:

$\frac{\partial}{\partial t}(\rho \vec{u})$ momentum change

$\vec{\nabla}\left[\rho\left(\underline{\underline{T}} + \vec{u} \cdot \vec{u}\right)\right]$ pressure force

$\rho \vec{\nabla} \Phi$ grav. force

Let's look for some familiar ground ...

If $\underline{\underline{T}}$ has the simple isotropic form

$$\underline{\underline{T}} = \begin{pmatrix} p & 0 & 0 \\ 0 & p & 0 \\ 0 & 0 & p \end{pmatrix}$$

as for an „ideal gas" and if the system is in steady state $\left(\vec{u} \equiv 0, \frac{\partial}{\partial t} \equiv 0 \right)$, then we get

$$\boxed{\vec{\nabla} p(\vec{x}) = -\rho(x) \vec{\nabla} \Phi(x)}$$

simple hydrostatic equilibrium

Before getting serious about solving the „Jeans Equation", let's play the integration trick one more time ...

Virial Theorem

Consider for simplicity the one-dimensional analog of the Jeans Equation in steady state:

$$\frac{\partial}{\partial x}\left[\rho v^2\right] + \rho\frac{\partial \Phi}{\partial x} = 0$$

After integrating over velocities, let's now

integrate over $\vec{x}$: $\int \ldots x d\vec{x}$

[one needs to use Gauss' theorem etc..]

$$\boxed{-2E_{kin} = E_{pot}}$$

Application of the Jeans Equation

- Goal:
  - Avoid "picking"right virial radius.
  - Account for spatial variations
  - Get more information than "total mass"

- Simplest case
- spherical: $\rho(\vec{r}) = \rho(r)$

  static: $\vec{\upsilon} \equiv 0, \frac{\partial}{\partial t} \equiv 0$

$$\vec{\nabla}\left(\rho \underline{\underline{T}}\right) = -\rho\vec{\nabla}\Phi$$

Choose spherical coordinates: $\frac{d}{dr}\left(\rho\sigma_r^2\right) + \frac{2\rho}{r}\left(\sigma_r^2 - \sigma_t^2\right) = -\rho\frac{d\Phi}{dr}$

σ_r is the radial and σ_t the tangential velocity dispersion

$$\frac{d}{dr}\left(\rho\sigma_r^2\right) = -\rho\frac{d\Phi}{dr}$$

for the „isotropic" case!

Note: Isotropy is a mathematical assumption here, **not** justified by physics!

Remember: ρ is the mass density of particles under consideration (e.g. stars), while Φ just describes the gravitational potential acting on them.

How are ρ and Φ related?

Two options:

1. $\nabla^2\Phi = 4\pi G\rho$ „self-consistent problem"
2. $\nabla^2\Phi = 4\pi G\underbrace{\left(\rho + \rho_{other}\right)}_{\rho_{total}}$ with

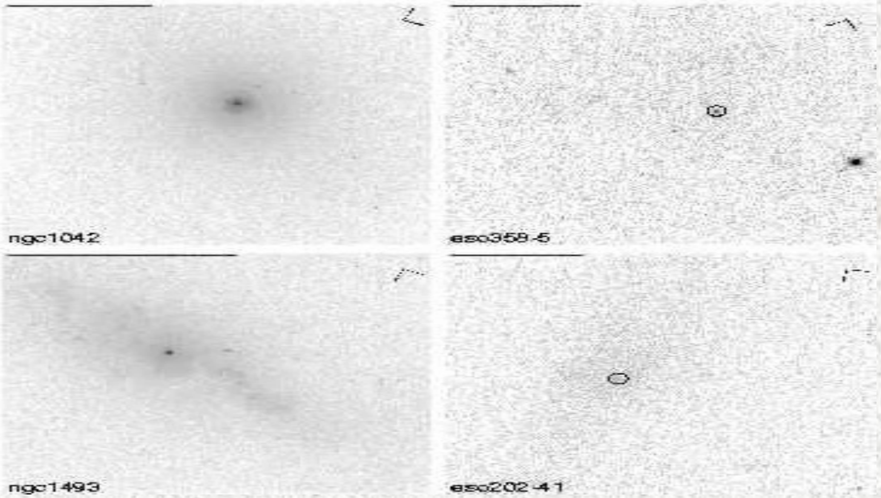
$\rho_{other} = \rho_{dark\ matter} + \rho_{gas} + \ldots \rho_{Black\ Hole}$

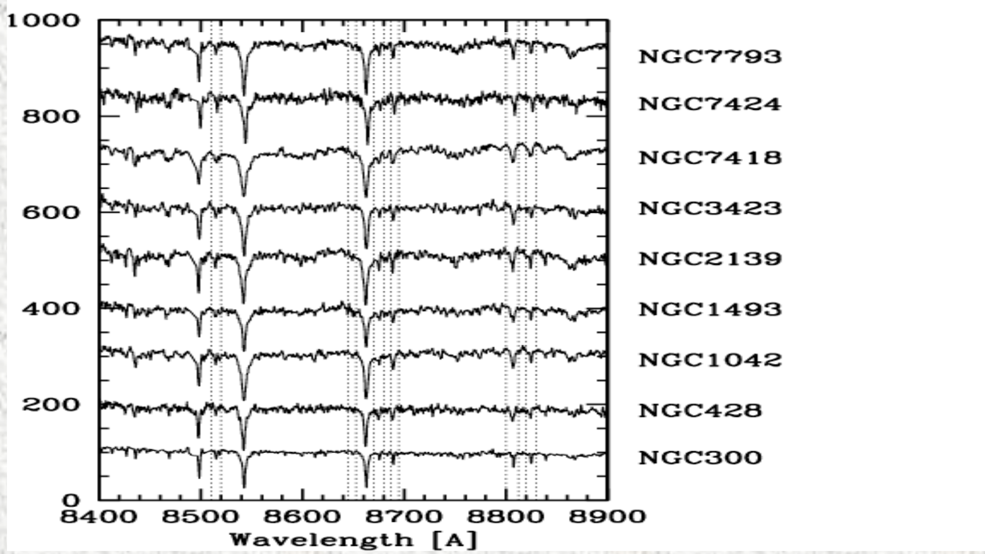
An Example:
When Jeans Equation Modeling is Good Enough

Walcher et al 2003, 2004

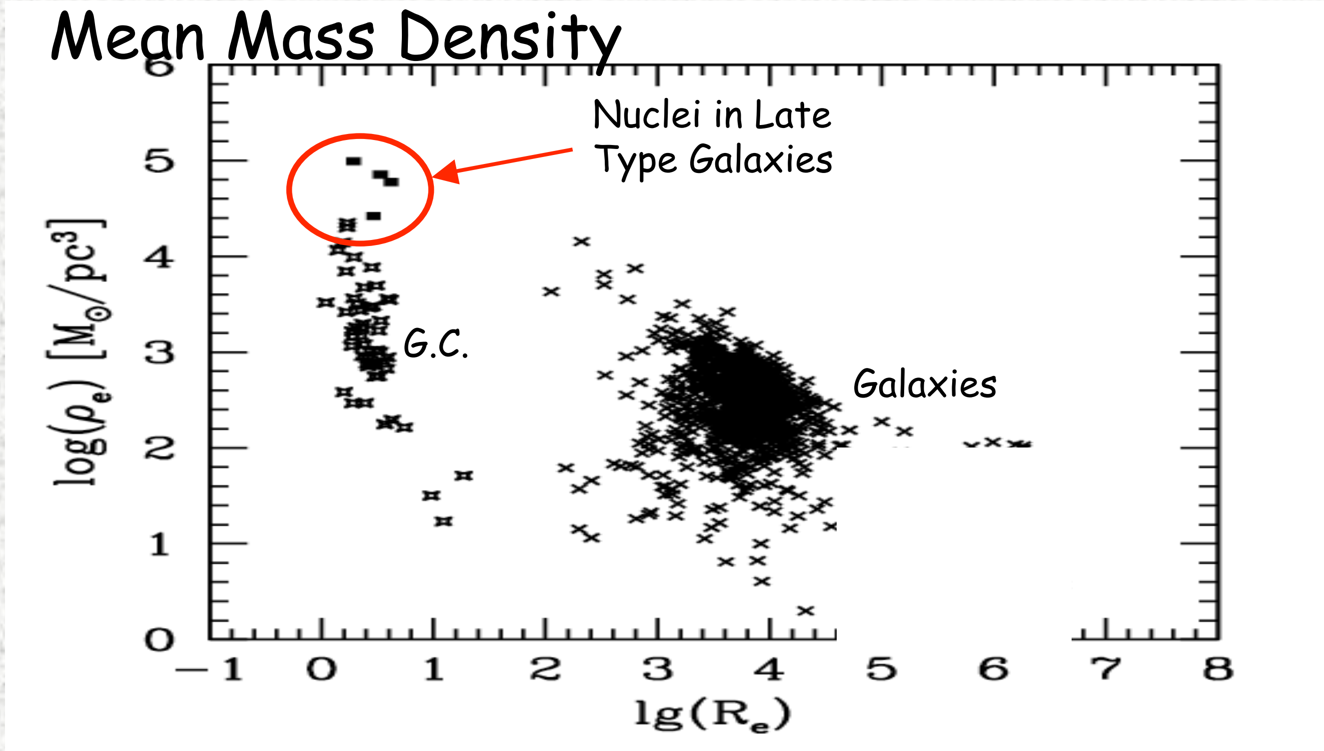
The densest stellar systems sitting in very diffuse galaxies..

Images → $\rho(r)$ Spectra → σ (perhaps $\sigma(r)$)





Then get M from the Jeans Equation ....



- Jeans Equation is great for estimating total masses for systems with limited kinematic data

Describing Collisionless Systems: Approach II

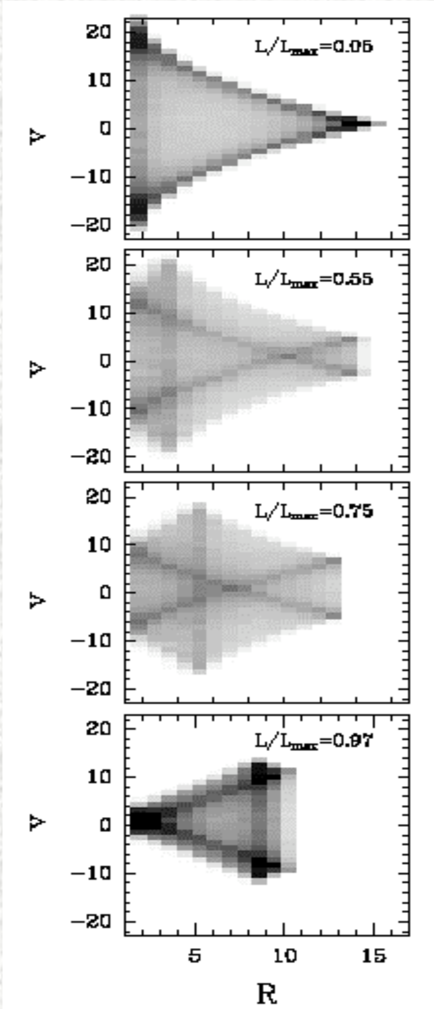
"Orbit-based" Models

Schwarzschild Models (1978)

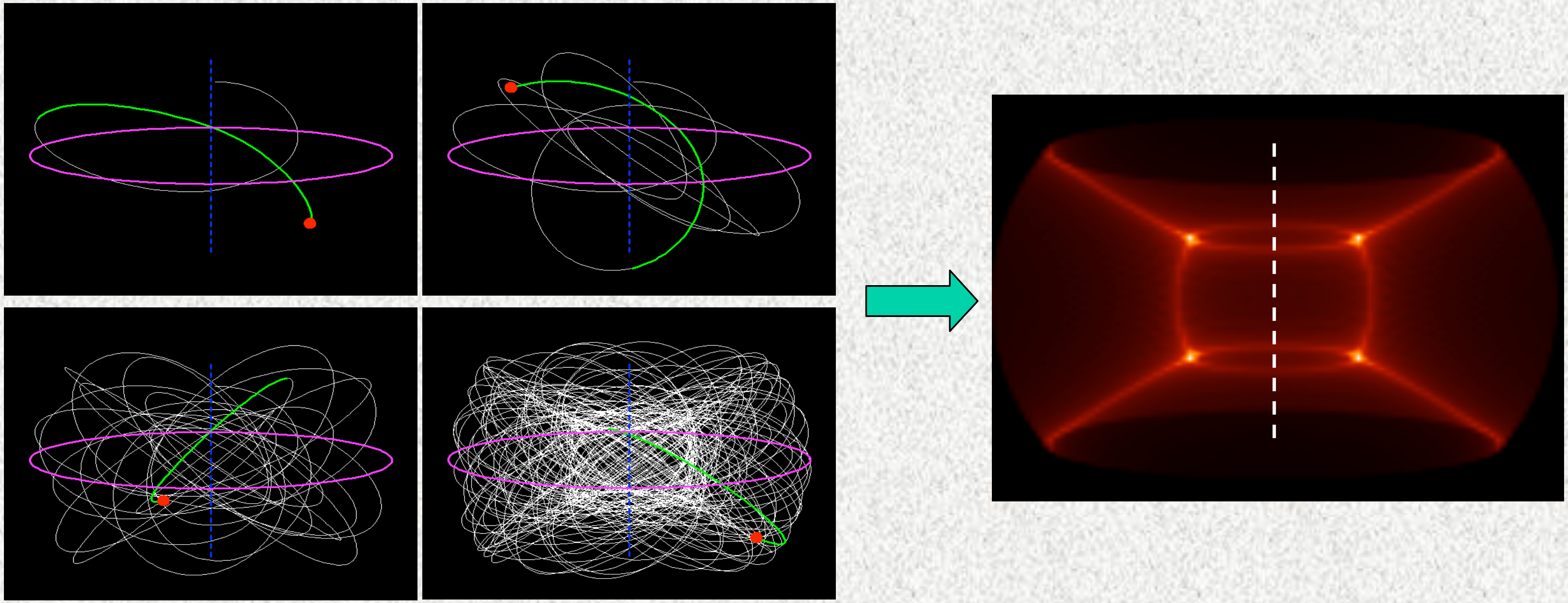
- What would the galaxy look like, if all stars were on the same orbit?
  - pick a potential Φ
  - Specify an orbit by its "isolating integrals of motion", e.g. E, J or J_z
  - Integrate orbit to calculate the
    - time-averaged
    - projected

    properties of this orbit

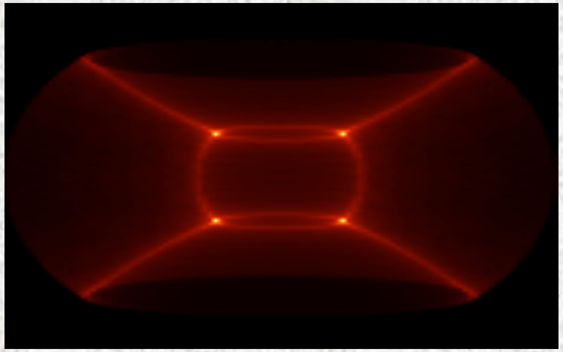
    (NB: time average in the calculation is identified with ensemble average in the galaxy at on instant)
  - Sample "orbit space" and repeat



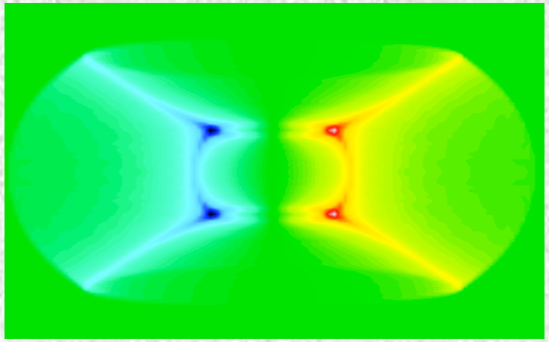
from Rix et al 1997



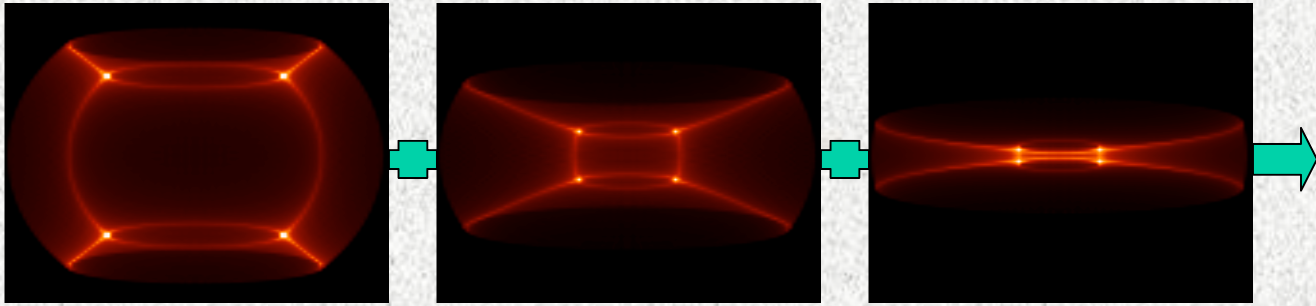
Figures courtesy Michele Capellari 2003



Projected density

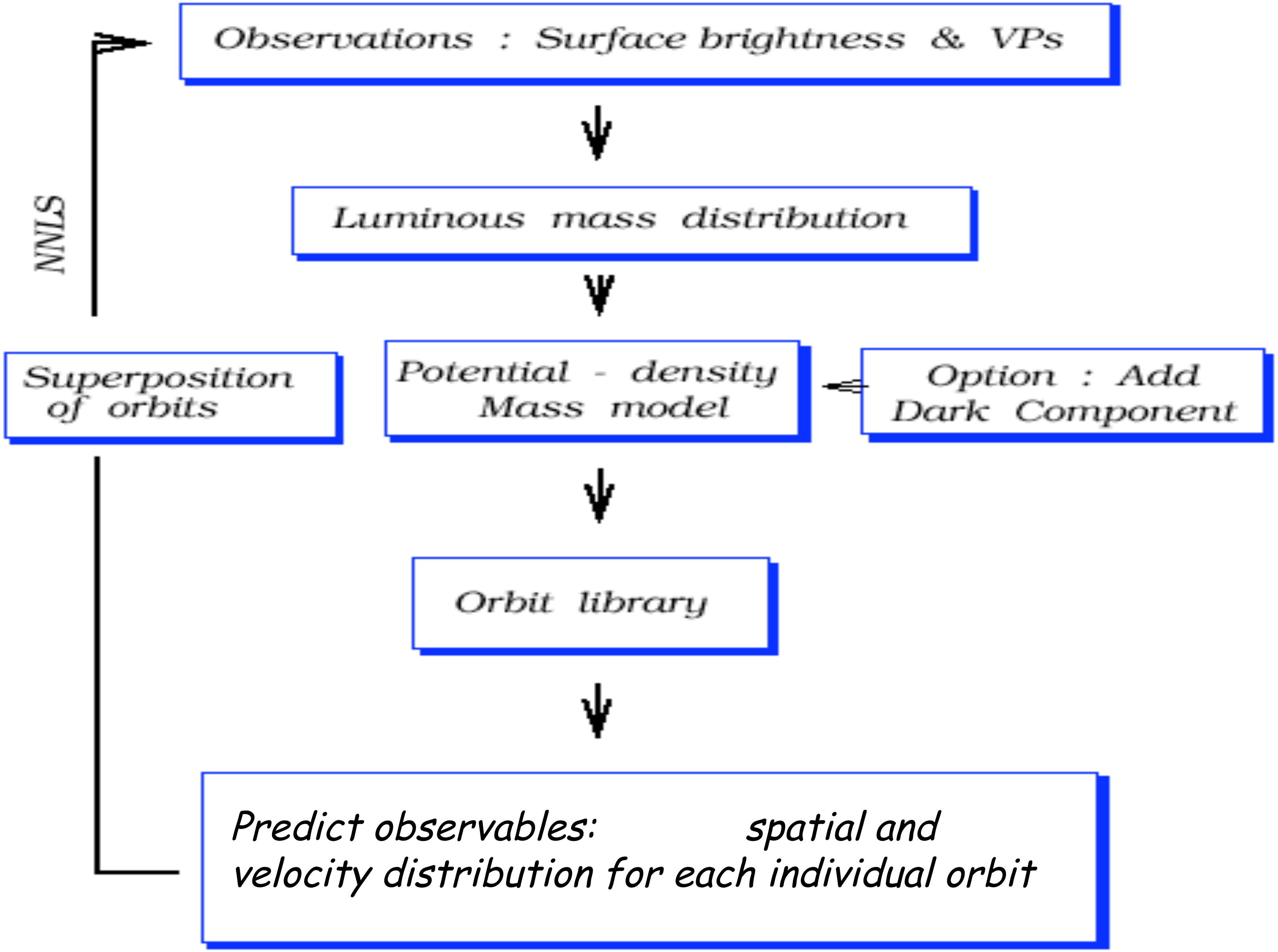


$V_{line\text{-}of\text{-}sight}$



images of model orbits

Observed galaxy image

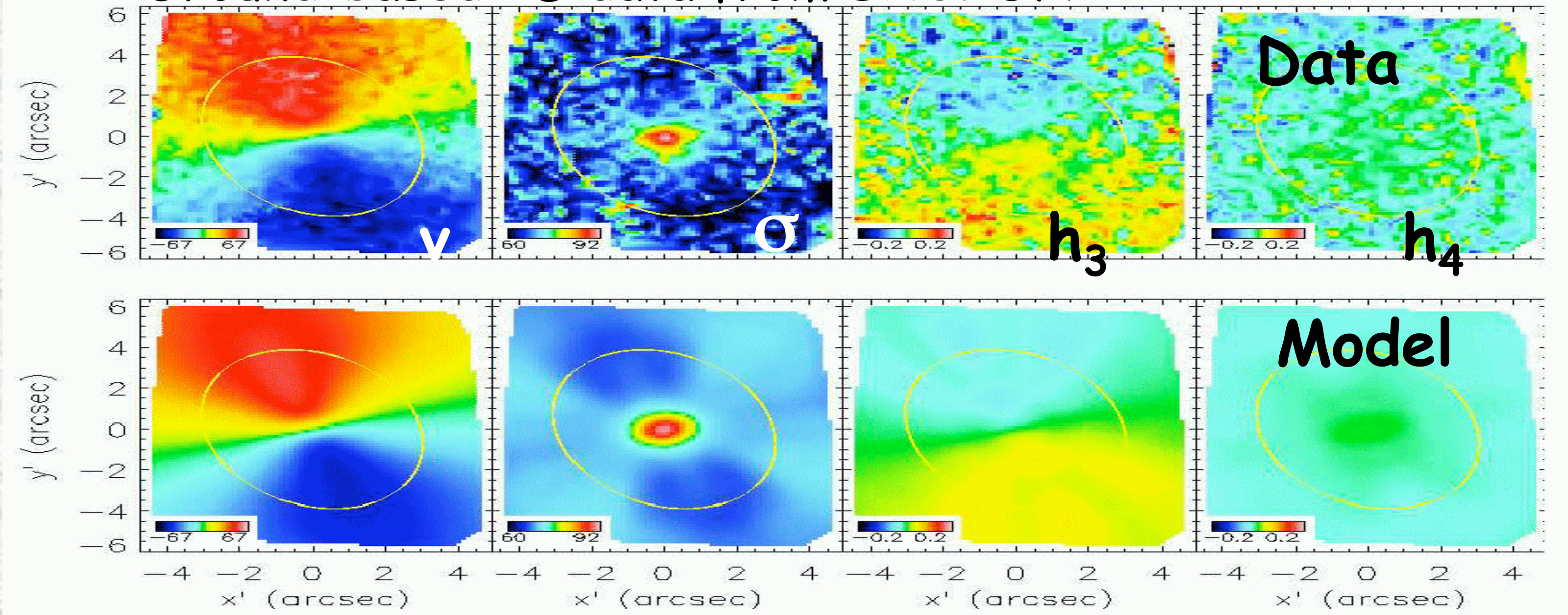

Observations : Surface brightness & VPs
Luminous mass distribution
Potential - density Mass model
Option : Add Dark Component
Superposition of orbits
NNLS
Orbit library
Predict observables: spatial and velocity distribution for each individual orbit

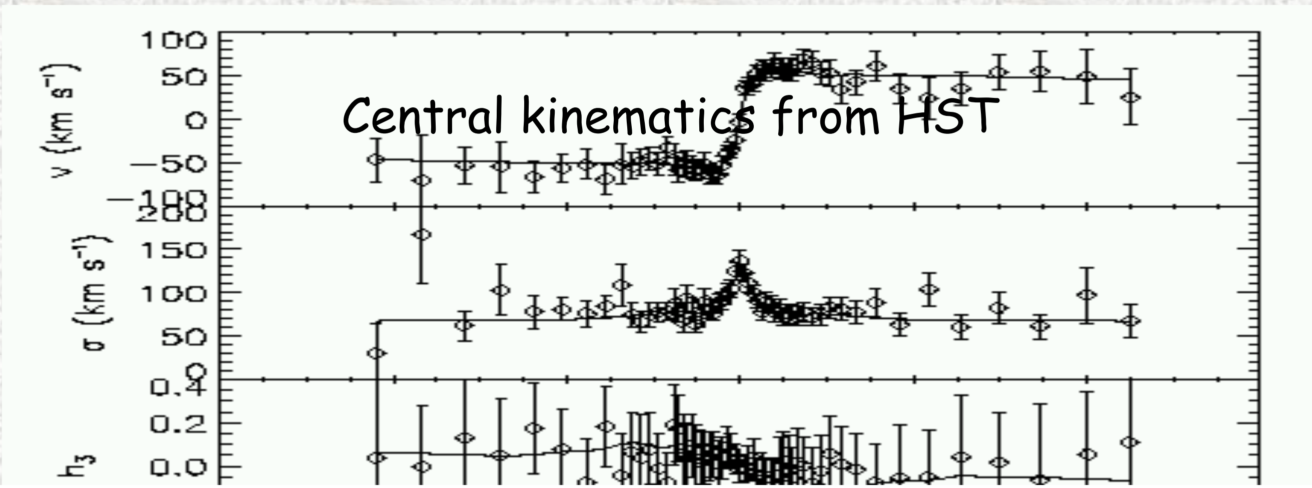
Example of Schwarzschild Modeling

M/L and M_{BH} in M32

Verolme et al 2001

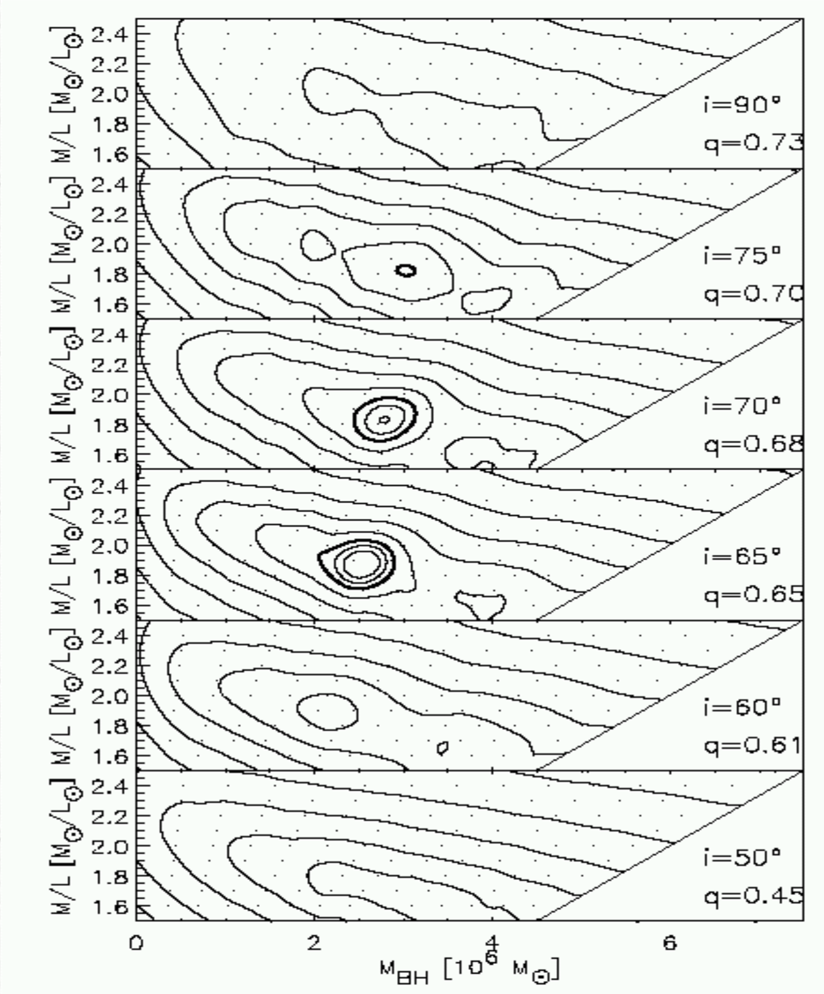
Ground-based 2D data from SAURON





- Then ask:

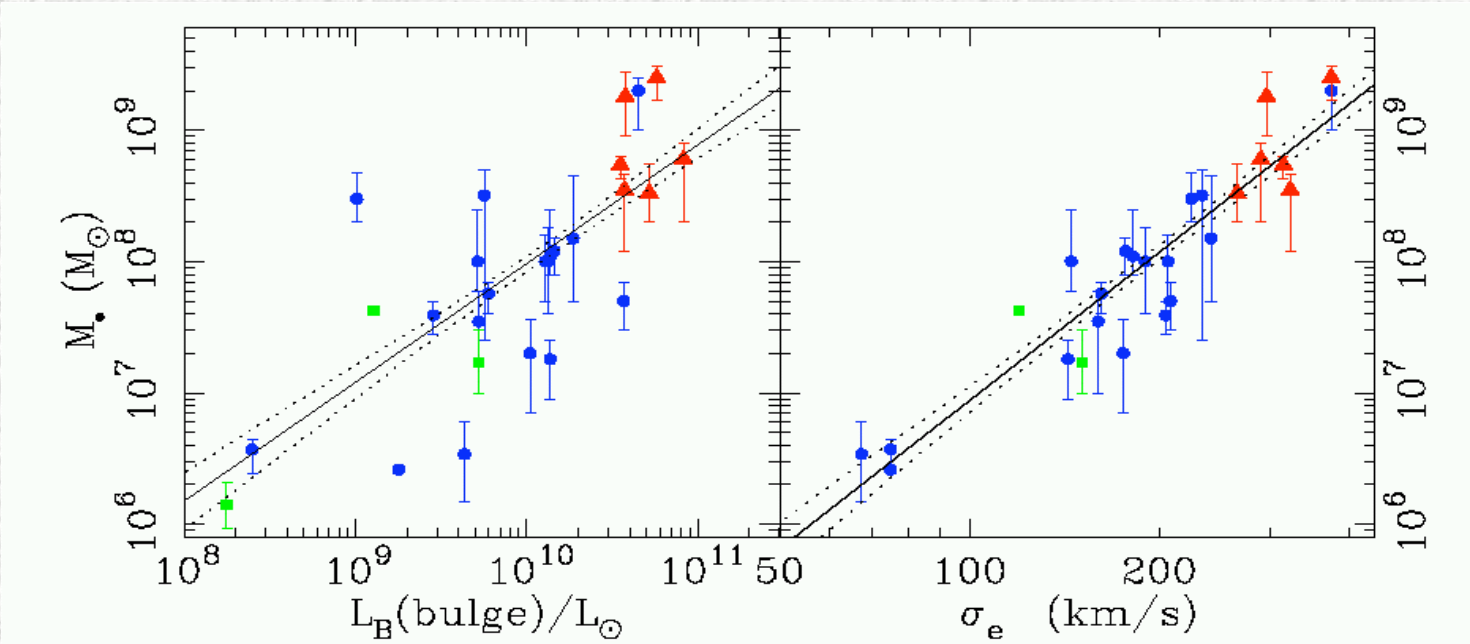
  for what potential and what orientation, is there a combination of orbits that matches the data well



Determine: inclination, M_{BH} and M/L simultaneously

NB: assumes axisymmetry

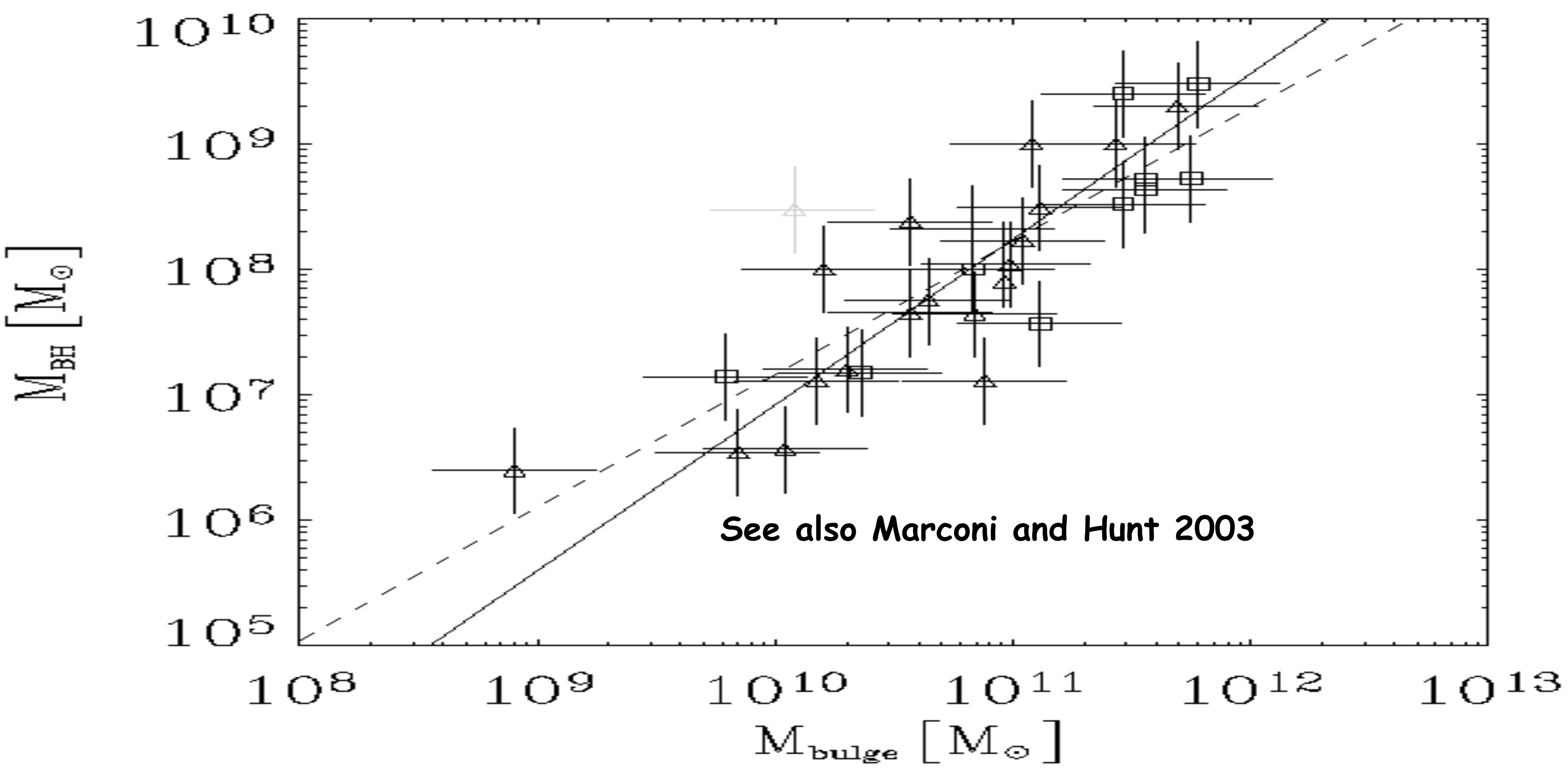
This type of modeling (+HST data) have proven necessary (and sufficient) to determine M_{BH} dynamically in samples of nearby massive galaxies



→ M_{BH} and σ_* (on kpc scales) are tightly linked (Gebhardt et al 2001)

B.t.w.: M_{BH} vs $M_{*,Bulge}$ seems just as good

Haering and Rix 2003



Stellar Kinematics and Clues to the Formation History of Galaxies

Mergers scramble the dynamical structure of galaxies, but do not erase the memory of the progenitor structures completely.

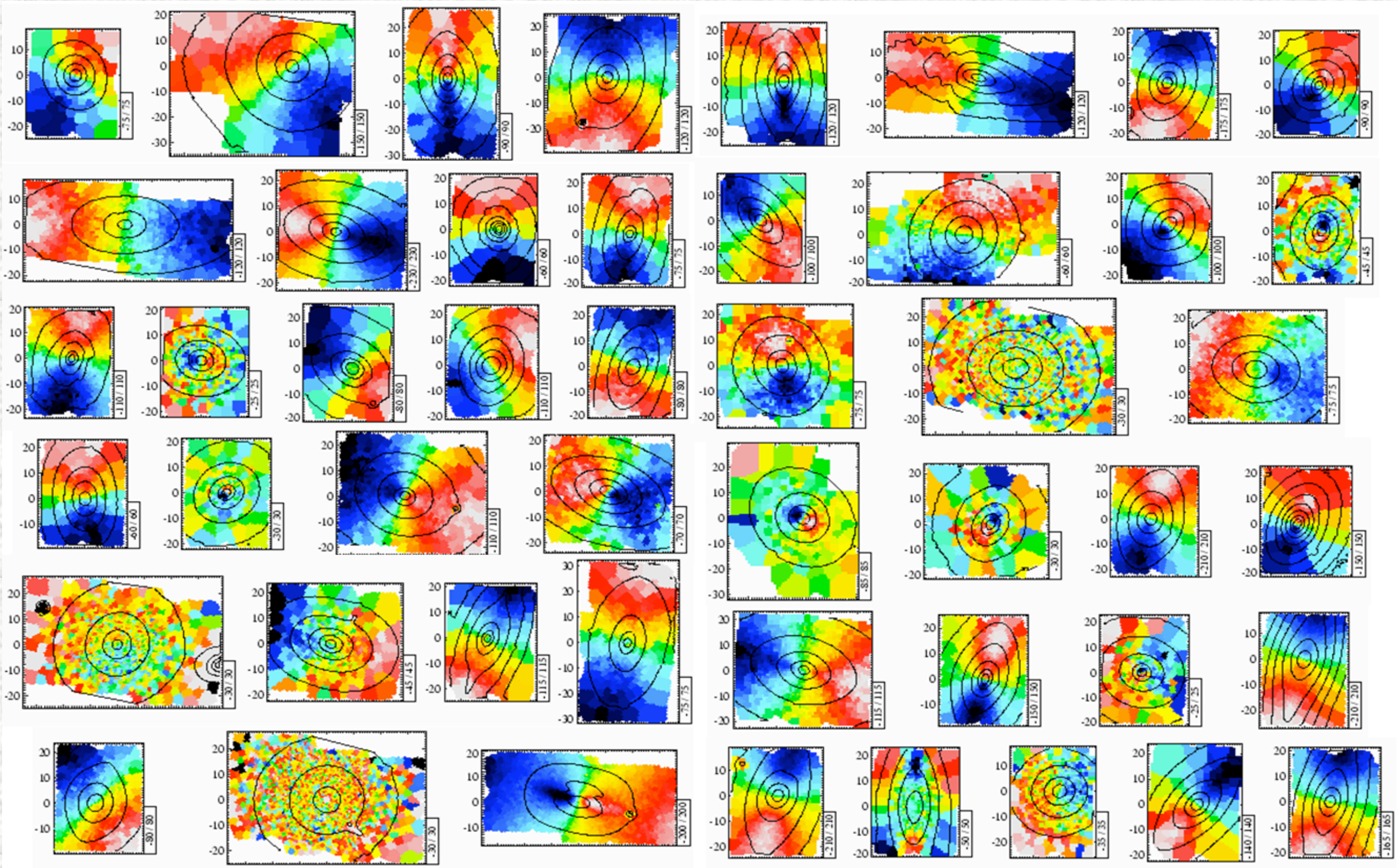
In equilibrium, phase space structure (E,J/Jz,+) is preserved.

However, observations are in $x_{proj}, y_{proj}, v_{LOS}$ space!

Connection not trivial!

Let's look at spheroids

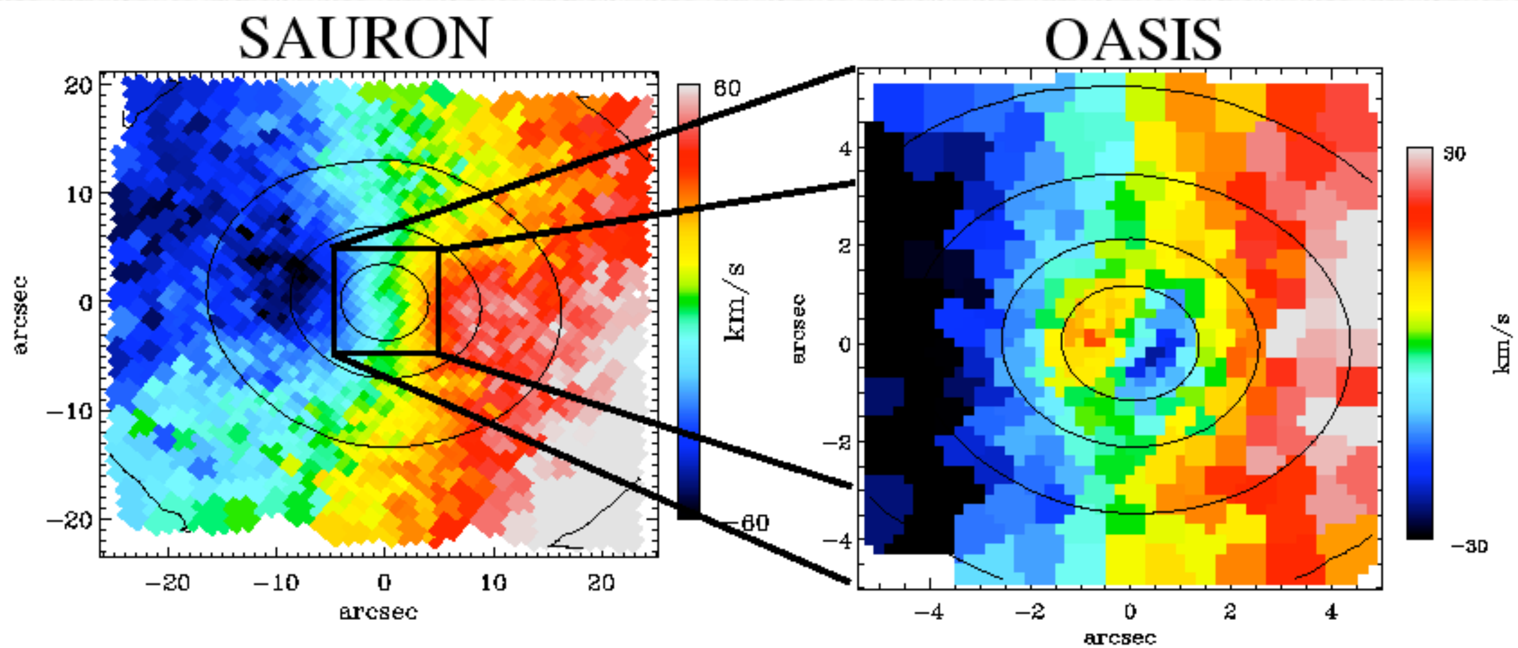
(data courtesy of the SAURON team
Co-PIs: de Zeeuw, Davies, Bacon)



Emsellem et al. 2003, MNRAS, submitted

SAURON versus OASIS

(total body vs. center)

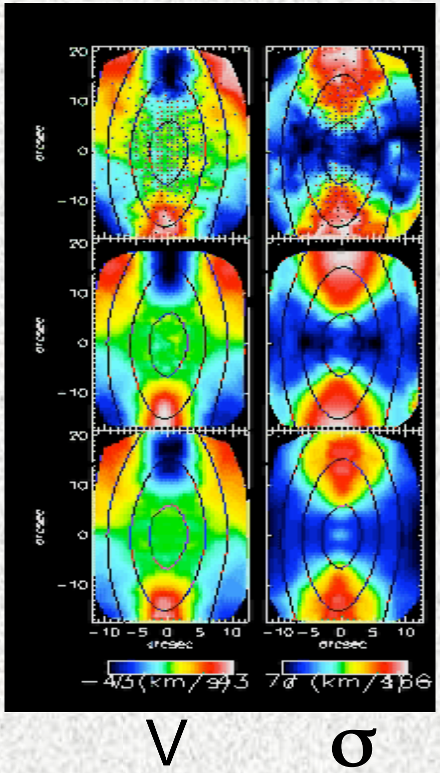


NGC 4382 McDermid et al. (2003) astro-ph/0311204

- Cores often have different (de-coupled) kinematics!

Intriguing Aside: NGC4550

a disk galaxy with ½ the stars going the wrong way?
(Rubin et al, Rix et al 1993)



2D-binned data

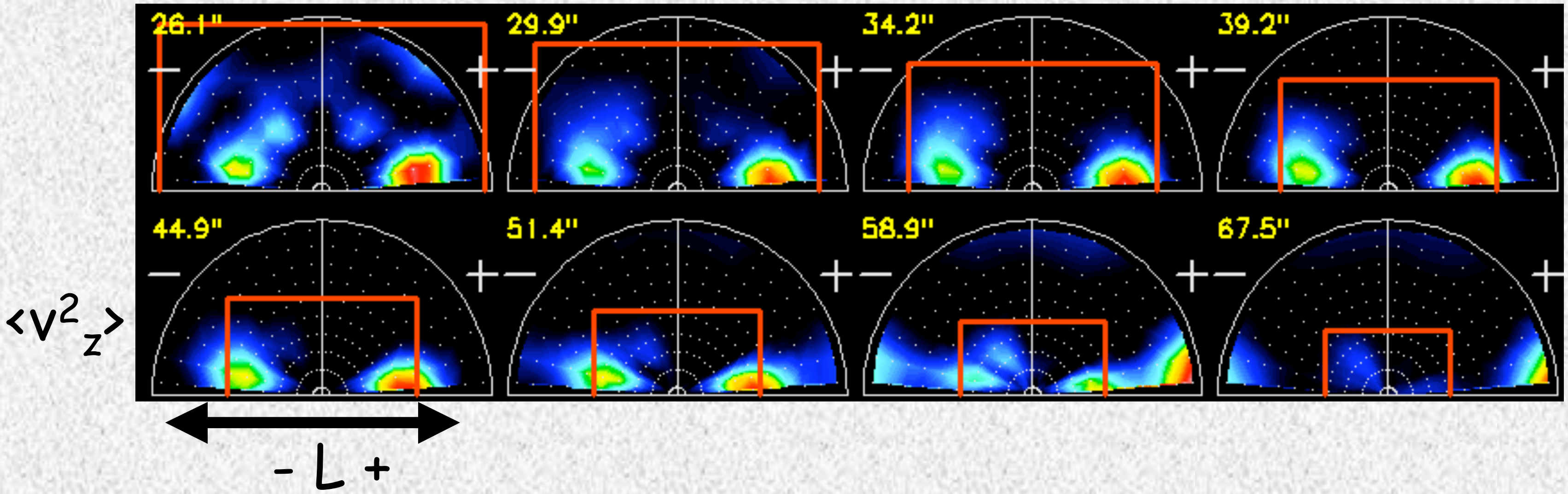
Symmetrized data

Axisymmetric model
M/L = 3.4 ± 0.2

- Axisymmetric dynamical model fits up to h5-h6
- M/L very accurate

NGC 4550: phase-space density

(that is solution of the Schwarzschild model)



- Two counterrotating components
  - Double-peaked absorption lines (Rix et al. 1993, ApJ, 400, L5)
  - SAURON: accurate decomposition, in phase space
- Both components are disks
  - Same mass
  - Different scale height

II. Some Basic Concepts in Non-Equilibrium Stellar Dynamics

a) Dynamical friction

b) Conservation of phase-space structure

c) Tidal disruption

d) Violent relaxation

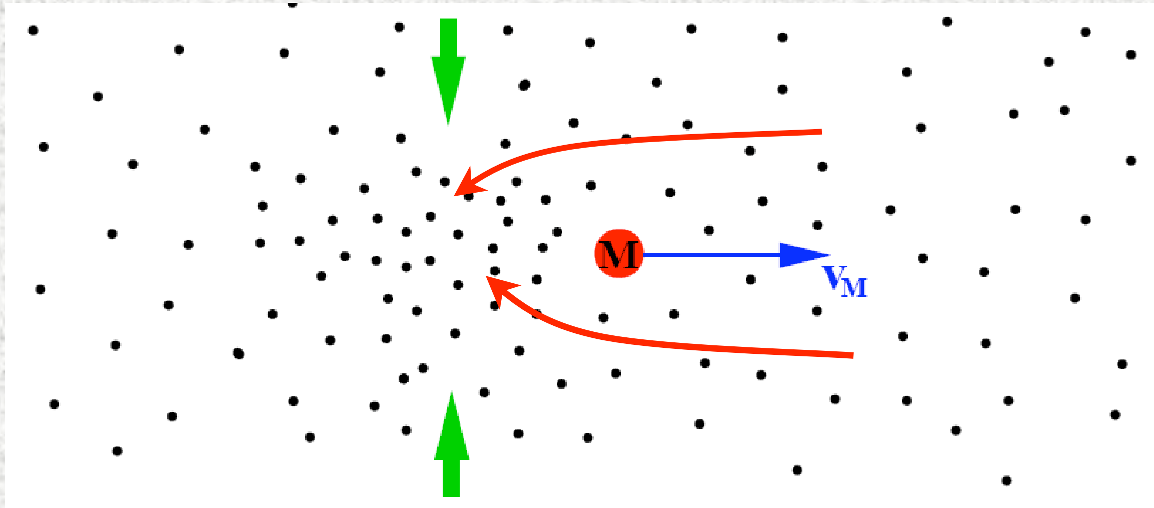
a) Dynamical friction

A "heavy" mass, a satellite galaxy or a bound sub-halo, will experience a slowing-down drag force (dynamical friction) when moving through a sea of lighter particles

Two ways to look at the phenomenon

a) A system of many particles is driven towards "equipartition", i.e. $E_{kin}(M) \sim E_{kin}(m)$
=> $V^2_{\text{of particle } M} < V^2_{\text{of particle } m}$

b) Heavy particles create a 'wake' behind them



$$F_{dyn.fric} = -\frac{4\pi G M^2}{V_M^2}\rho_m \cdot \ln\Lambda$$

Where m << M and ρ_m is the (uniform) density of light particles m, and $\Lambda = b_{max}/b_{min}$ with $b_{min} \sim \rho_M/V_2$ and $b_{max} \sim$ size of system typically lu $\Lambda \sim 10$

Effects of dynamical Friction

a) Orbital decay: $t_{df} \sim r\ /\ (dr/dt)$

V_{circ} dr/dt = -0.4 ln Λ ρ_M/r

or

$$t_{df} \approx \frac{1.2}{lu\Lambda}\frac{r_i^2 V_c}{\rho M}$$

Example: orbital decay of the Large Magellanic Cloud in MW Halo

$V_{cir}(MW) = 220$ km/s $M_{LMC} = 2 \times 10^{10}\ M_{SUN}$ $R_i = 50$ kpc

$\Rightarrow$ $t_{df} = 1.2$ Gyr

b) Galaxy Merging

Ultimately, galaxy (or halo) merging is related process. The (heavy) bound part of one merger participant is transferring its orbital energy to the individual (light) particles of the other merger participant (and vice versa).

Issues:

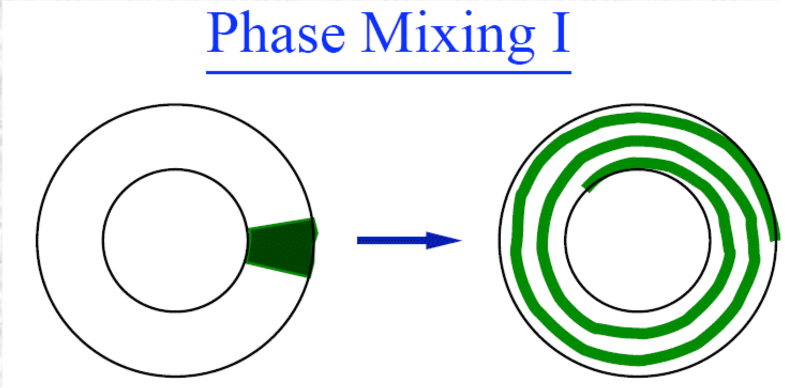
1. Merging preserves ordering in binding energy (i.e. gradients)
2. Merging destroys disks – isotropizes
3. 'Dry' merging (i.e. no gas inflow) lowers (phase-space) density
4. Post-merger phase-mixing makes merger looks smooth in ~few t_{dyn}

b) Conservation of phase-space structure

In stationary, or slowly-varying potentials:

- Sub-structure is phase-mixed in 'real space'
- Phase-space density (e.g. E,L_z space) is conserved

→ dynamics basis of 'galactic archeology'



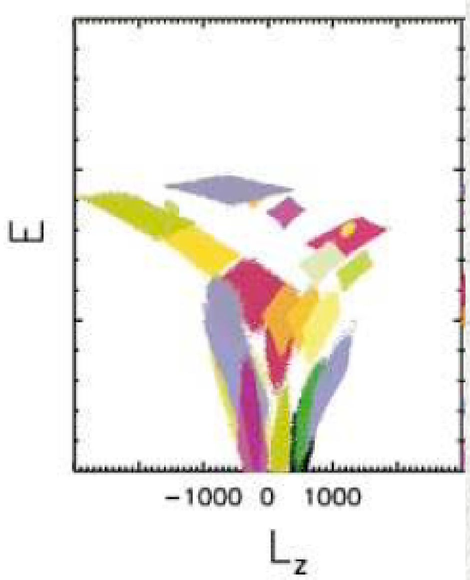
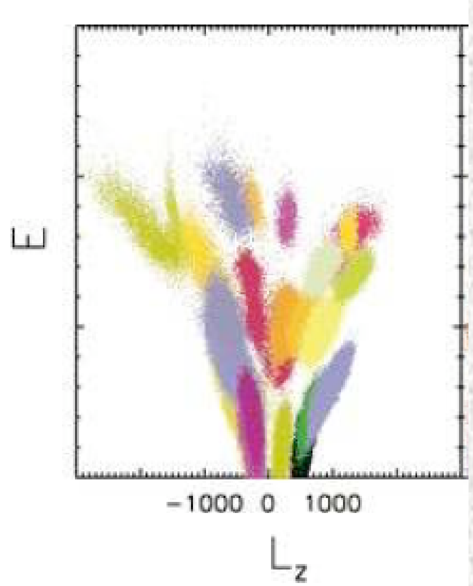


Figure 3. Initial distribution of particl

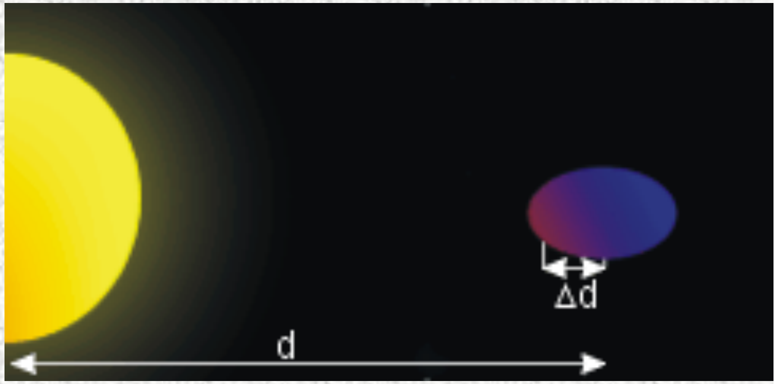
Initial clumps in phase-space →. observed 10Gyrs later with tha GAIA satellite



c) Tidal disruption

‘ "Roche limit": for existence of a satellite, its self-gravity has to exceed the tidal force from the ‘parent’

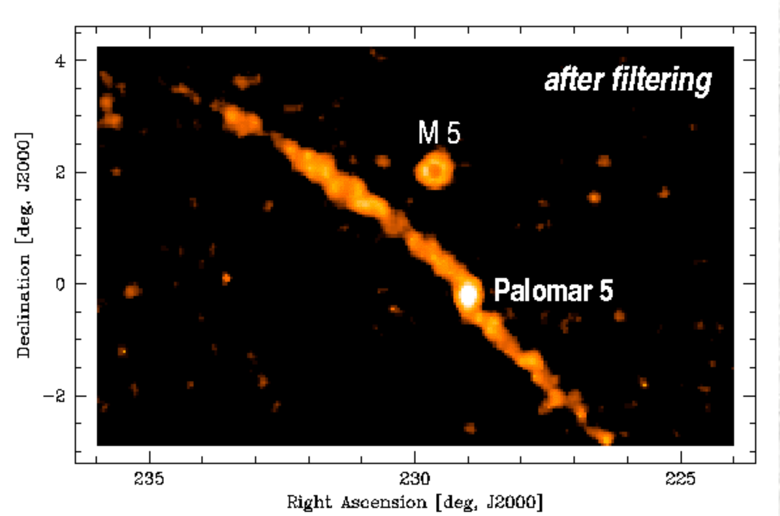
$$V_T = -\frac{3GM}{2d^3}\Delta d^2$$



Tidal radius: $$R_{tidal}(satellite) = f\left[\frac{M_{satellite}}{M_{host}(< R_{peri})}\right]^{1/3} \times R_{peri} \quad \text{with } f \approx 2/3[1-\ln(1-e)]^{-1/3}$$

In cosmological simulations, many DM sub-halos get tidally disrupted.

- How important is it, e.g. in the Milky Way?
- The GC Pal 5 and the Sagitarius dwarf galaxy show that it happens



d) Violent relaxation

Mon. Not. R. astr. Soc. (1967) **136**, 101–121.

STATISTICAL MECHANICS OF VIOLENT RELAXATION IN STELLAR SYSTEMS

D. Lynden-Bell

(Communicated by the Astronomer Royal)

(Received 1966 December 19)

Summary

An explanation of the observed light distributions of elliptical galaxies is sought and found.

The violently changing gravitational field of a newly formed galaxy is effective in changing the statistics of stellar orbits.

Basic idea:

- (rapidly) time-varying potential changes energies of particles
- Different change for different particles

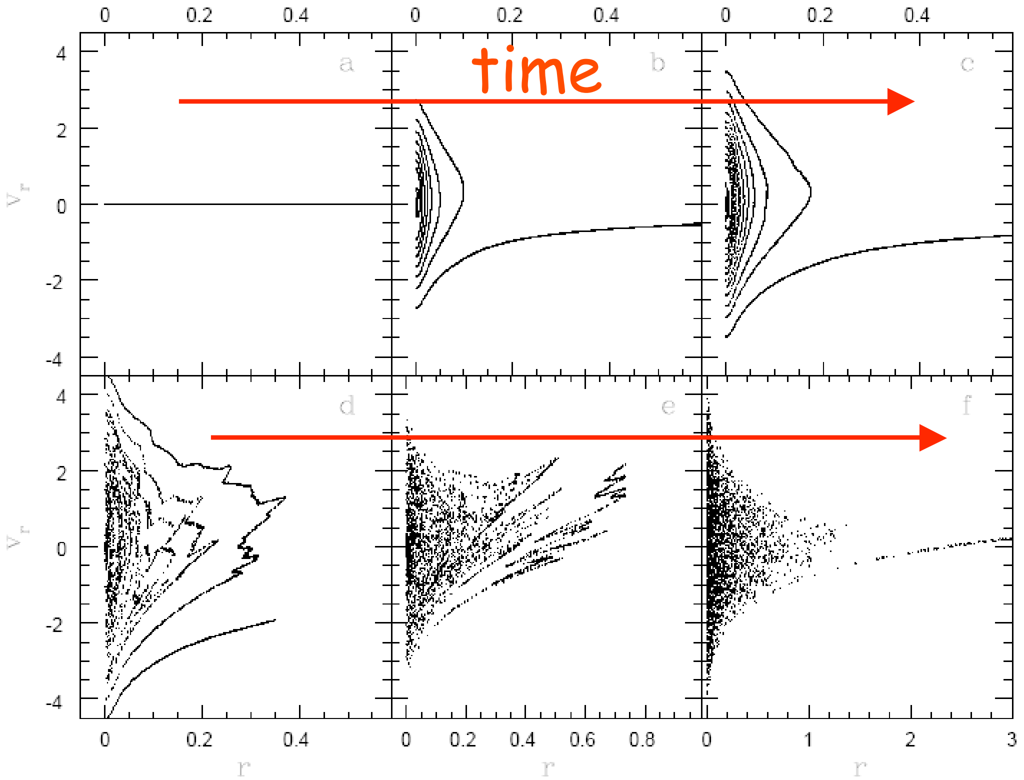
$$E = \frac{1}{2}v^2 + \Phi \text{ and } \Phi = \Phi(\vec{x}, t)$$

$$\frac{\mathrm{d}E}{\mathrm{d}t} = \frac{\partial E}{\partial \vec{v}} \cdot \frac{\mathrm{d}\vec{v}}{\mathrm{d}t} + \frac{\partial E}{\partial \Phi}\frac{\mathrm{d}\Phi}{\mathrm{d}t} = \frac{\partial \Phi}{\partial t}$$

The time-scale for violent relaxation is

$$t_{\rm vr} = \left\langle \frac{({\rm d}E/{\rm d}t)^2}{E^2} \right\rangle^{-1/2} = \left\langle \frac{(\partial\Phi/\partial t)^2}{E^2} \right\rangle^{-1/2} = \frac{3}{4}\langle \dot{\Phi}^2/\Phi^2 \rangle^{-1/2}$$

How violent relaxation works in practice (i.e. on a computer)



(from: Henriksen & Widrow 1997)

Collapse of a spherical system with $\rho_{\rm init} \propto r^{-3/2}$

III. The Dynamics of Gas in Galaxies (vs. Stars)

Two regimes:

- $KT \approx V^2_{characteristic}$ hot gas
- $KT \ll V^2_{characteristic}$ warm, cold gas

Dynamics of 'hot' gas

'approximate hydrostatic equilibrium'
X-ray gas, 10^6 K observable in massive galaxies

b) Dynamics of 'cold' gas

To 'avoid' shocks, gas will settle on non-intersecting loop orbits:

•concentric circles (in axisymm. case)

•ellipses in (slightly distorted) potentials

•E.g. weak spiral arms

•in barred potentials, closed-orbit ellipticity changes at resonances → shocks, inflow

•Observed:

•Bars drive gas inflow

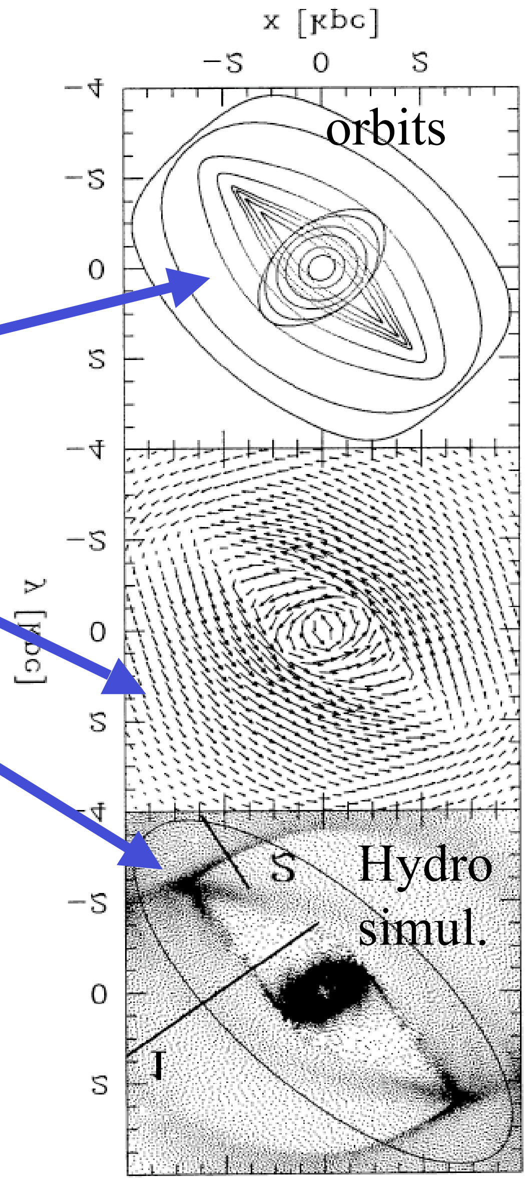
(e.g. Schinnerer et al 07)

•Whether all the way to

the black hole, unclear..

Gas Flow in Non-Axisymmetric Potential

(Englmaier et al 97)



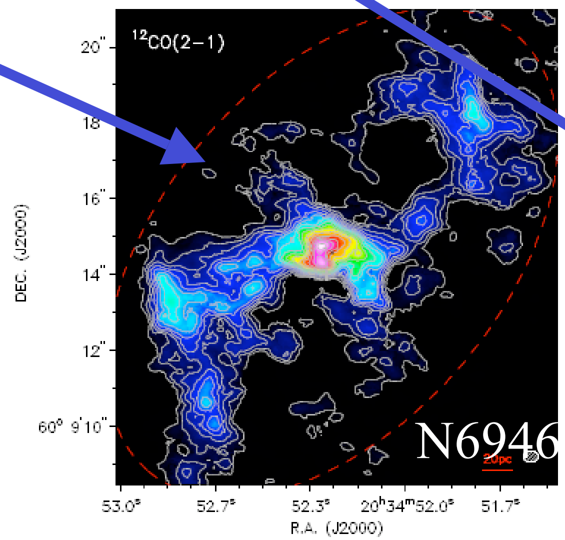


Fig. 1. Intensity map of the ^{12}CO(2-1) line emission at 0.4"×0.3" resolution (color and contours). The black cross marks the posi-

Dynamics Summary

- Collisionless stars/DM and (cold) gas have different dynamics
- "Dynamical modeling" of equilibria
  - Answering: In what potential and on what orbits to tracers move?
  - Two approaches: Jeans Equation vs. Orbit (Schwarzschild) modeling
  - N.B: 'kinematic tracers' need not cause the gravitational potential
    most modeling assumes random orbital phases; not true if $t_{orb} \sim t_{Hubble}$
- Phase-space density (e.g. in E, L, L_z coordinates) conserved in static or slowly varying potentials → dynamical archeology?
  - 'violent relaxation' may erase much of this memory
- (Cold) gas dynamics : dissipational not collisionless matter
  - Wants to form disks
  - In (strongly) non-axisymmetric potentials: shocks → inflow
  - No phase-space 'memory'